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Interacting Dark Sector (ETHOS $n = 0$): Cosmological Constraints from SPT Cluster Abundance with DES and HST Weak Lensing Data

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We use galaxy cluster abundance measurements from the South Pole Telescope (SPT) enhanced by Multi-Component Matched Filter (MCMF) confirmation and complemented with mass information obtained using weak-lensing data from Dark Energy Survey Year 3 (DES Y3) and targeted Hubble Space Telescope (HST) observations for probing deviations from the cold dark matter paradigm. Concretely, we consider a class of dark sector models featuring interactions between dark matter (DM) and a dark radiation (DR) component within the framework of the Effective Theory of Structure Formation (ETHOS). We focus on scenarios that lead to power suppression over a wide range of scales, and thus can be tested with data sensitive to large scales, as realized for example for DM–DR interactions following from an unbroken non-Abelian $SU(N)$ gauge theory (interaction rate with power-law index $n = 0$ within the ETHOS parameterization). Cluster abundance measurements are mostly sensitive to the amount of DR interacting with DM, parameterized by the ratio of DR temperature to the cosmic microwave background (CMB) temperature, $\xi_{\text{DR}} = T_{\text{DR}}/T_{\text{CMB}}$. We find an upper limit $\xi_{\text{DR}} < 17\%$ at 95% credibility. When the cluster data are combined with Planck 2018 CMB data along with baryon acoustic oscillation (BAO) measurements we find $\xi_{\text{DR}} < 10\%$, corresponding to a limit on the abundance of interacting DR that is around three times tighter than that from CMB+BAO data alone. We also discuss the complementarity of weak lensing informed cluster abundance studies with probes sensitive to smaller scales, explore the impact on our analysis of massive neutrinos, and comment on a slight preference for the presence of a non-zero interacting DR abundance, which enables a physical solution to the S_8 tension.

I. INTRODUCTION

The abundance of galaxy clusters has long been recognized as a powerful cosmological probe [1, 2], and in recent years, cluster number counts have provided competitive constraints on cosmological parameters [3–6]. Clusters are the largest collapsed structures in the Universe,

and their abundance is particularly sensitive to the density of matter Ω_{m} , the amplitude of density fluctuations at large scales, usually parameterized by σ_8 at $8h^{-1}$ Mpc, and the dark energy equation of state parameter w . This makes cluster number counts an important tool to test models that have an impact on structure formation and deviate from the standard cosmological model (Λ CDM).

Understanding the nature of dark matter (DM) is a central problem of cosmology. No satisfactory hint for the identity of DM beyond its gravitational interaction has so

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far been detected. Observations of the cosmic microwave background anisotropies (CMB) [7] and galaxy clustering via baryon acoustic oscillations (BAO) [8–10], together with weak lensing information [11–13], can be used as precision probes of certain fundamental DM properties. Future CMB and large scale structure (LSS) surveys [14–20] are expected to be sensitive to even small deviations from Λ CDM. In that regard, the tension in the value of $S_8 = \sigma_8 \sqrt{\Omega_m}/0.3$, can serve as a reference for searches beyond the cold dark matter (CDM) scenario. For Λ CDM, a tension of $\sim 1.5 - 3\sigma$, depending on the dataset considered, has been reported between the high values obtained by the Planck 2018 analysis of CMB anisotropies [7] and structure formation probes such as KiDS [11], DES [12], HSC [21], SPT [4], eROSITA (eFEDS) [22], eBOSS [23], except eROSITA (eRASS1) [6] which reported a higher value than CMB. Even in the absence of any strong evidence of deviations from the CDM paradigm, quantifying the extent to which the behaviour of DM may deviate from that of a cold and collisionless matter component is still highly informative for identifying viable DM theories.

Extensions of DM models beyond the CDM paradigm include for example DM self-interaction on galactic [24, 25] or even cosmological scales [26, 27] and decays to other types of particles, see e.g. [28–30]. In this work we focus on the commonly considered scenario of a secluded dark sector where DM (or a fraction of it) interacts with a relativistic component called dark radiation (DR) [31–43]. In this case, structure formation can be (slightly) suppressed on scales $\gtrsim h^{-1}$ Mpc, making these types of models relevant for addressing the S_8 tension. Interacting dark matter–dark radiation (IDM–DR) scenarios arise within many classes of theoretical models, and their cosmological effect can be described within the framework of the Effective Theory of Structure Formation (ETHOS) [34, 35]. A comparison between different models and their impact on structure formation can be found in [20].

A particularly straightforward and predictive example for a microscopic IDM–DR model is based on a dark sector featuring a weakly coupled, unbroken $SU(N)$ gauge symmetry [32, 33, 39, 43]. In this scenario, DM consists of a particle species with gauge interactions governed by this symmetry, and the dark gauge bosons constitute the DR component. Their characteristic self-coupling gives rise to an interaction between IDM and DR that leads to a moderate suppression of the matter power spectrum over a wide range of scales [32], such that this setup can explain low S_8 values and be tested with observations sensitive to relatively large scales [20, 39, 41, 43, 44]. Given that $SU(N)$ gauge theories underlie both the strong and weak force in Nature, it is a plausible question whether interactions of DM are governed by a similar mechanism. Note that, cosmologically, DM–DR interactions within a dark sector are analogous to baryon–photon interactions within the visible sector. The IDM–DR scenario thus probes a well-

motivated, fundamental property of DM, being described by a framework that is conceptually closely related to the baryon–photon drag force underlying CMB acoustic oscillations.

In previous analyses [41, 43] this model has been explored using Planck measurements [7] of CMB anisotropies, as well as BAO and galaxy clustering data from BOSS DR12 [8]. Results indicate a preferred region in parameter space where σ_8 (or S_8) drops to lower values in the presence of DR. In a recent analysis [44], the sensitivity of galaxy cluster counts to this model was investigated, including a forecast for ongoing and future surveys like SPT-3G and CMB-S4 complemented with next-generation weak-lensing data, such as those expected from the Euclid and Rubin surveys. This analysis found that cluster counts are particularly sensitive to the DR energy density, or equivalently the value of the temperature ratio $\xi_{\text{DR}} = T_{\text{DR}}/T_{\text{CMB}}$, and will be able to discriminate between Λ CDM and IDM–DR models. A forecast of the sensitivity of weak-lensing shear measurements by Euclid [20] as well as for future 21cm observations [45] has been carried out within a slightly different setup of the same model.

In this work, we perform an analysis of the IDM–DR model based on the abundance of galaxy clusters detected via the thermal Sunyaev-Zeldovich effect (tSZE) in the South Pole Telescope (SPT) [46] survey data and confirmed using the MCMF algorithm [47, 48] applied to optical/NIR followup data. This cluster abundance analysis is informed by weak lensing data from the Dark Energy Survey (DES) Y3 [49–51] and targeted weak-lensing measurements from the Hubble Space Telescope (HST) [52]. The cluster sample comprises 1,005 clusters constructed from the combined SPT-SZ, SPTpol ECS, and SPTpol 500d surveys [53–56], in the redshift range 0.25 – 1.78. Weak-lensing data provide information on cluster masses and allow for constraining observable–mass relations empirically. We employ weak-lensing data from DES Year 3 for 688 clusters with redshift $z < 0.95$ and HST data for 39 clusters with a redshift range 0.6–1.7. The same combination of datasets has recently been analyzed in [5] and shown to yield competitive constraints within Λ CDM as well as for the dark energy equation of state and the sum of neutrino masses. We follow the same data analysis framework in this work.

The structure of this article is as follows: in Sec. II, we review the IDM–DR setup and summarize the SPT cluster and HST/DES weak-lensing datasets and analysis strategy in Sec. III and Sec. IV, respectively. Our main results are discussed in Sec V, and we conclude in Sec. VI. The Appendices contain results for an extended model setup as well as posteriors including all parameters entering the joint cluster abundance and weak-lensing mass calibration analysis.

II. IDM–DR MODEL SETUP

Interacting dark sector models can be realized by various particle physics setups. Yet, structure formation on large scales is only sensitive to certain characteristics of DM. The ETHOS framework offers a systematic approach for describing these models in terms of a few relevant parameters that encapsulate the impact on structure formation, and for mapping them to the underlying particle physics model properties [34]. ETHOS thus allows for constraining a wide class of generic DM models using observational cosmological data in an efficient way. In this framework, one can differentiate between various models based on the temperature (and hence redshift) dependence of the effective interaction rate between DM and DR. Specifically, the rate relevant for the drag force generated by DM–DR interactions within the dark sector can be parameterized as $\Gamma_{\text{IDM-DR}}(z) \propto \sum_n a_n (1+z)^{n+1}$. Here n characterizes the dependence on redshift z , and a_n is the absolute interaction strength of inverse length dimension. Both of these parameters encapsulate the properties of the microscopic model. In previous works [41, 43], analyses were carried out assuming for simplicity an interaction rate described by a single power-law, and considering the cases $n = 0, 2, 4$, that are characteristic for various classes of particle models differing by the dependence of the DM–DR cross section on the momentum transfer. The case $n = 0$ turns out to be phenomenologically most interesting as it gives rise to an interaction rate that features the same redshift-dependence as the Hubble rate during radiation domination. This means the ratio $\Gamma_{\text{IDM-DR}}/\mathcal{H}$ (where \mathcal{H} is the conformal Hubble rate) stays constant over an extended period of time, impacting a wide range of perturbation modes with different scales entering the horizon, leading to a rather gradual suppression of the matter power spectrum, in sharp contrast to $n > 0$ for which a cutoff similar to warm dark matter scenarios is predicted [32]. In contrast, the power suppression extending over a wide range of scales obtained from IDM–DR interactions with $n = 0$ offers a potential solution to the S_8 tension [20, 39, 41, 43, 44]. Furthermore, the $n = 0$ case is particularly well-motivated from a particle physics perspective. Concretely, the dark sector interaction described by a weakly coupled, unbroken non-Abelian $SU(N)$ gauge theory predicts an IDM–DR interaction strength described by $n = 0$ within ETHOS. In what follows, we will describe its main properties and quantities of interest.

Following previous work, we allow for the possibility that only a fraction of the total DM population interacts with DR, given by

$$f_{\text{IDM}} \equiv \frac{\Omega_{\text{IDM}}}{\Omega_{\text{IDM}} + \Omega_{\text{CDM}}}, \quad (1)$$

with Ω_{IDM} and Ω_{CDM} being the IDM and CDM density parameters. Within the non-Abelian dark sector model, the IDM component is described by a Fermionic particle

species that transforms under a non-trivial representation of $SU(N)$, assuming the fundamental representation for concreteness. The non-interacting part of the DM is provided by a particle species that transforms trivially under $SU(N)$. The DR consists of a thermal bath of massless $SU(N)$ gauge bosons. Their self-interactions, being a characteristic feature of a non-Abelian gauge group, imply that DR behaves as a *fluid* component (rather than as free-streaming radiation). In general, the dark sector is allowed to have a different temperature than the visible sector, with a ratio that remains constant throughout the cosmological epochs considered in this work, and given by

$$\xi_{\text{DR}} \equiv \frac{T_{\text{DR}}}{T_{\text{CMB}}}. \quad (2)$$

The two parameters f_{IDM} and ξ_{DR} are the relevant model parameters for our analysis in the following. For convenience, we also present various derived parameters, such as the DR density parameter

$$\Omega_{\text{DR}} = (N^2 - 1) \xi_{\text{DR}}^4 \Omega_{\gamma}, \quad (3)$$

in terms of Ω_{γ} , the photon density parameter today. The contribution of DR to the effective number of relativistic species is given by

$$\Delta N_{\text{eff}} = \frac{\rho_{\text{DR}}}{\rho_{1\nu}} = \frac{8}{7} \left(\frac{11}{4} \right)^{4/3} (N^2 - 1) \xi_{\text{DR}}^4, \quad (4)$$

with ρ_{DR} and $\rho_{1\nu}$ being, respectively, the energy densities for DR and one massless neutrino family.

The interaction between IDM and DR affects the evolution of density and velocity perturbations in a way that is analogous to the well-known baryon–photon interactions. The set of evolution equations for the density contrast δ_{IDM} and velocity divergence θ_{IDM} obtained within the ETHOS framework is given by

$$\dot{\delta}_{\text{IDM}} + \theta_{\text{IDM}} - 3\dot{\phi} = 0, \quad (5)$$

$$\dot{\theta}_{\text{IDM}} - c_{\text{IDM}}^2 k^2 \delta_{\text{IDM}} + \mathcal{H} \theta_{\text{IDM}} - k^2 \psi = \Gamma_{\text{IDM-DR}} (\theta_{\text{IDM}} - \theta_{\text{DR}}), \quad (6)$$

where $k = |\mathbf{k}|$ is the comoving wave number and c_{IDM} is the adiabatic IDM sound speed. The gravitational potentials ϕ and ψ within conformal Newtonian gauge act as a source for perturbations for continuity and Euler's equations, as usual. Importantly, within ETHOS, the Euler equation features a drag term depending on the relative difference between θ_{IDM} and the DR velocity divergence θ_{DR} . This interaction rate can be written as a power law in redshift as

$$\Gamma_{\text{IDM-DR}}(z) = -\frac{4}{3} (\Omega_{\text{DR}} h^2) a_{\text{dark}} (1+z),$$

with a_{dark} a free parameter for quantifying the interaction strength, that has units of Mpc^{-1} . Here h is the

dimensionless Hubble constant. Note that this redshift-dependence matches the general parameterization discussed above, with power-law index $n = 0$. The DR component is treated as a fluid, in line with the non-Abelian self-interaction of dark gauge bosons, and therefore we only take into account the first two moments of its distribution: the overdensity δ_{DR} and the velocity divergence θ_{DR} , for which we can write

$$\dot{\delta}_{\text{DR}} + \frac{4}{3}\theta_{\text{DR}} - 4\dot{\phi} = 0, \quad (7)$$

$$\begin{aligned} \dot{\theta}_{\text{DR}} - \frac{1}{4}k^2\delta_{\text{DR}} \\ + k^2\sigma_{\text{DR}}^2 - k^2\psi = \Gamma_{\text{DR-IDM}}(\theta_{\text{DR}} - \theta_{\text{IDM}}), \end{aligned} \quad (8)$$

with shear stress σ_{DR} being set to zero in the fluid limit, and interaction rate $\Gamma_{\text{IDM-DR}} = (4\rho_{\text{DR}}/3\rho_{\text{IDM}}) \times \Gamma_{\text{DR-IDM}}$ fixed by energy-momentum conservation with ρ_{IDM} and ρ_{DR} being the IDM and DR energy densities.

While we stress that the setup presented here applies in principle to all IDM-DR dark sector models with interaction rate that has a redshift-dependence described by $n = 0$ within ETHOS, we note that within the dark sector described by the $SU(N)$ gauge interaction the IDM-DR rate can be related to its fundamental parameters, which include the DM mass m_χ and the dark gauge coupling constant g_d [43],

$$\begin{aligned} \Gamma_{\text{IDM-DR}}(z) = -\frac{1}{1+z} \frac{\pi}{18} \frac{\alpha_d^2}{m_\chi} 2(N^2 - 1) \times \\ \left\{ T_{\text{DR}}^2 [\ln \alpha_d^{-1} + c_0 + c_1 g_d + \mathcal{O}(g_d^2)] + \mathcal{O}\left(\frac{T_{\text{DR}}^4}{m_\chi^2}\right) \right\}, \end{aligned} \quad (9)$$

where $\alpha_d = g_d^2/(4\pi)$ is the dark fine structure constant, and

$$c_0 = 1 + \ln\left(\frac{6}{2N + N_f}\right) + \ln(4\pi) - 24 \ln(A), \quad (10)$$

$$c_1 = \frac{3\sqrt{2N + N_f}}{4\pi} \sqrt{\frac{3}{2}}, \quad (11)$$

with the Glaisher-Kinkelin constant $A \simeq 1.28243$ and N_f being the number of additional light Fermionic degrees of freedom in the dark sector (with $N_f = 0$ in the most minimal version of the model). Eq. (9) takes into account the Debye-screening by the DR plasma (see [43] for a derivation). Using this result, the interaction strength parameter a_{dark} can be mapped to the fundamental parameters,

$$\begin{aligned} a_{\text{dark}} &= \frac{\pi}{12} \frac{\alpha_d^2}{m_\chi} \frac{(N^2 - 1)^{\frac{1}{2}} T_{\text{CMB},0}^2}{\xi_{\text{DR}}^2 \Omega_\gamma h^2} \\ &\times [\ln \alpha_d^{-1} + c_0 + c_1 g_d + \mathcal{O}(g_d^2)] \\ &= 0.91 \cdot 10^9 \text{Mpc}^{-1} \left(\frac{\alpha_d}{10^{-4}}\right)^2 \left(\frac{100 \text{ GeV}}{m_\chi}\right) \left(\frac{0.1}{\xi_{\text{DR}}}\right)^2 \\ &\times (N^2 - 1)^{\frac{1}{2}} [\ln \alpha_d^{-1} - (1.34 + \ln N) + 0.413\sqrt{N}g_d]. \end{aligned} \quad (12)$$

In this work we focus on the tight coupling limit, where the value of a_{dark} is large enough to keep the IDM and DR tightly coupled. This requires $a_{\text{dark}} \gtrsim 10^6 \text{Mpc}^{-1}$ (see [44] for more details). Using Eq. (12), we stress that the tight coupling limit is easily realized, requiring only a very weak lower bound on the dark fine-structure of $\alpha_d \gtrsim 10^{-6}$, for $m_\chi = 100 \text{ GeV}$ and $\xi_{\text{DR}} = 0.1$ assuming an $SU(3)$ theory. To further constrain the value of α_d , we use other properties of the model. In a non-Abelian $SU(N)$ gauge theory the interaction becomes stronger at low energies because of the running coupling, which can lead to confinement. To ensure that the cosmological evolution occurs within the unconfined phase, we require $\alpha_d \lesssim 0.1$. Furthermore, the observed ellipticity of the gravitational potential of the galaxy NGC720 leads to a bound on the long-range interaction strength [57], being $\alpha_d \lesssim 5 \cdot 10^{-3}$ for $m_\chi = 100 \text{ GeV}$ and $N = 3$. More discussion on both constraints can be found in [43]. Finally, IDM self-interaction with a cross-section of order $1 \text{ cm}^2/\text{g}$ is relevant for small-scale puzzles, like the core-cusp problem (see [24] for a review and also [31, 35]). This can be achieved by a dark fine structure constant of order $\alpha_d \sim 10^{-5}$, depending on m_χ . We note that this value is consistent with the lower bound required for the tight coupling limit as well as the upper bounds from confinement and galaxy ellipticity.

In summary, this IDM-DR setup can be characterized by three parameters: the fraction f_{IDM} of interacting DM, Eq. (1), the amplitude a_{dark} of this interaction, Eq. (12), and the temperature of dark radiation relative to the CMB ξ_{DR} , Eq. (2). Increasing a_{dark} and f_{IDM} enhances the suppression of the matter power spectrum relative to ΛCDM , but this has only a minor impact on the halo mass function within the regime probed by cluster counts, resulting in limited constraining power from this type of observation. In contrast, varying ξ_{DR} leads to a change of the scale on which the matter power spectrum is suppressed, having a substantial impact on the halo mass function, and consequently leading to a high sensitivity of galaxy cluster abundance measurements to this model parameter. For a more complete description of the effect of each parameter, we refer to [44]. Throughout the main part of this work we consider the tightly coupled regime of large a_{dark} , and therefore in our analysis the two free model parameters, ξ_{DR} and f_{IDM} , are sufficient to characterize the IDM-DR model.

III. CLUSTER CATALOG AND WEAK-LENSING DATASETS

We summarize the datasets used for the galaxy cluster abundance analysis in this work, covering first the MCMF enhanced SPT cluster catalog and then the DES Y3 and HST weak-lensing data, which are used for cluster mass calibration. Further details can be found in [58].

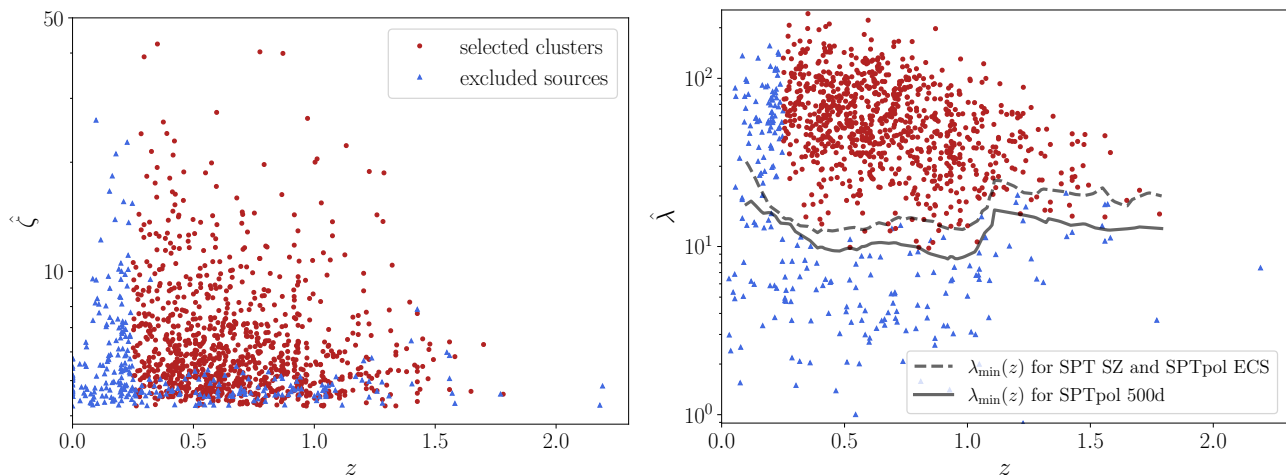


FIG. 1. The distribution of clusters and contaminating sources from the SPT surveys [54–56] in the observables tSZE significance $\hat{\zeta}$, optical or NIR richness $\hat{\lambda}$ and redshift z . The red points are confirmed clusters that meet the selection criteria listed in Eq. (13) and constitute the sample analysed here. The blue points at low $\hat{\zeta}$ (left) and $\hat{\lambda}$ (right) consist of contaminants in the original tSZE-selected candidate list and confirmed clusters at $z < 0.25$. The optical followup with the MCMF algorithm in cluster redshifts, excludes the contaminants and enhances the total number of confirmed clusters by $\sim 30\%$ in comparison to a purely tSZE selected sample with the same final contamination level of $\sim 2\%$. The figures illustrate that the bulk of the confirmed clusters lie well above the $\lambda_{\min}(z)$ MCMF selection threshold and that the bulk of the confirmed sample lies at redshifts $z < 1$. Note that the feature in the $\lambda_{\min}(z)$ selection at $z = 1.1$ corresponds to the inclusion of WISE data at higher redshift.

A. SPT cluster catalog

The SPT cluster catalog is constructed from three multi-frequency mm-wave surveys: SPT-SZ, SPTpol ECS, and SPTpol 500d [53–56]. Together, the surveys cover $5,270 \text{ deg}^2$ of the southern sky. Within the SPT-SZ observed area, the SPTpol 500d patch was re-observed to a greater depth, and we use only data from the latter in this overlapping region. Cluster candidates are selected using the observed tSZE *detection significance* $\hat{\zeta}$, which is the maximum detection signal-to-noise ratio (SNR) when scanning over a range of cluster angular sizes. Clusters are then confirmed using optical and NIR data, which delivers 1) the observed *richness*, $\hat{\lambda}$, corresponding to the weighted number of passive galaxies in a cluster, and 2) the cluster redshift z . For the clusters from the SPT-SZ and SPTpol surveys, the final sample size was increased by about $\sim 50\%$ using the MCMF algorithm, which allows one to use the distribution of richness measured along random lines of sight to push to lower tSZE detection significance $\hat{\zeta}$, while maintaining high purity in the final, confirmed cluster sample [47, 59]. This same technique has been employed to construct the largest all-sky X-ray selected cluster sample [60], the deepest Planck tSZE selected sample [61] and the largest tSZE selected cluster sample to date, which is based on ACT observations [62].

The cluster catalog over the entire survey region includes only clusters at $z > 0.25$. Clusters at lower redshift are excluded because the detection filter used for the SPT maps that removes atmospheric noise and noise contributions from the primary CMB measurement also

removes a significant fraction of the signal from the lower redshift clusters.

The overlapping region between the SPT surveys and DES Y3 weak-lensing dataset covers $3,567 \text{ deg}^2$ ($\sim 70\%$ of the SPT survey area). The confirmation process with MCMF [47] leads to the cluster redshift z , a richness $\hat{\lambda}$, and the optical center of the cluster on the sky. A tSZE selected SPT cluster candidate is considered a confirmed cluster only if its richness exceeds a richness threshold $\lambda_{\min}(z)$. To ensure the same sample purity (of $\sim 98\%$) at all redshifts, the lower limit in richness varies with cluster redshift $\lambda_{\min}(z)$; more details are provided in the SPT catalog papers [59, 63]. MCMF was run using optical DES Y3 galaxies as well as using infrared WISE galaxies [64]. At high redshifts ($z > 1$) the optical confirmation significantly suffers from the limited depth of DES Y3 data and is surpassed in sensitivity by the infrared confirmation at $z > 1.1$.

The various SPT surveys have different depths, resulting in cluster candidate lists with varying purity levels at a fixed detection significance, $\hat{\zeta}$. To maintain roughly constant purity in the cluster samples from the different SPT surveys, we apply different detection significance and richness thresholds. We select clusters based on the following criteria

$$\begin{aligned} \hat{\zeta} &> 4.25 / 4.5 / 5 \text{ (500d / SZ / ECS)}, \\ \hat{\lambda} &> \hat{\lambda}_{\min}(z), \\ z &> 0.25. \end{aligned} \quad (13)$$

For the $1,327 \text{ deg}^2$ of the SPT survey not covered by DES Y3, we select samples by applying a cut in the SPT

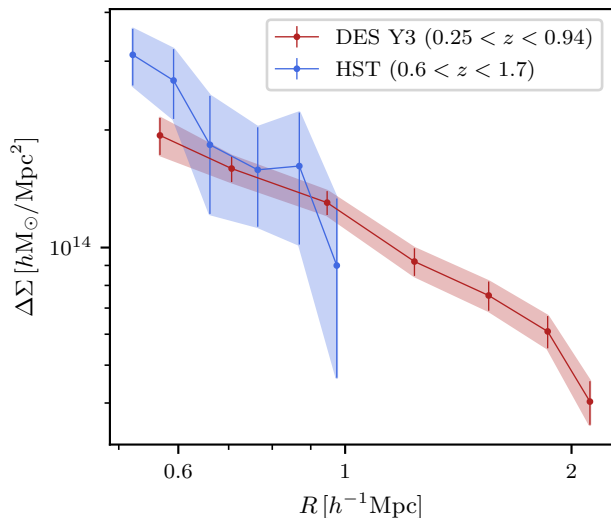


FIG. 2. Average matter profiles $\Delta\Sigma(R)$ for MCMF confirmed SPT selected clusters as extracted from DES Y3 data for 688 clusters (red) and from HST data for 39 clusters (blue). Error bars correspond to the 1σ region.

detection significance $\hat{\zeta} > 5$. The resulting cluster candidate list has a purity of $\gtrsim 95\%$. Cluster confirmation and redshift assignment are carried out using targeted optical observations (like PISCO; [65]). More details can be found in the original catalog publication [54].

The final sample consists of 1,005 confirmed clusters with tSZE significance, richness, and redshift measurements. Additionally, all clusters in the DES Y3 region have richness measurements and optically determined cluster centers provided with the MCMF algorithm operating on the DES Y3 or WISE datasets. We show the sample tSZE significance and richness as a function redshift in Fig. 1.

B. DES Y3 weak-lensing data

The DES Y3 dataset includes three years of survey data. It covers approximately $5,000 \text{ deg}^2$ of the southern sky. The data were collected using the 570-megapixel Dark Energy Camera (DECam) mounted on the Blanco 4-meter telescope at the Cerro Tololo Inter-American Observatory (CTIO) in Chile. The shape catalog was constructed using the METACALIBRATION method [66, 67]. Further information about systematics and calibration can be found in the DES Y3 catalog dedicated papers; [68] for photometric dataset, [69] for point-spread function modeling, and [70, 71] for image and survey simulations. After applying all source galaxy selection cuts, the DES Y3 shear catalog includes approximately 100 million galaxies over a solid angle of $4,143 \text{ deg}^2$ with an effective galaxy source density of $5\text{-}6 \text{ arcmin}^{-2}$.

A total of 3,567 deg^2 of the DES Y3 dataset over-

laps the SPT surveys. We select lensing source galaxies in four tomographic bins following the 3×2 pt analysis [12] with their redshift distribution calibrated using self-organizing maps [72]. For every SPT cluster, we extract a weak-lensing shear profile within the radial range $500 h^{-1} \text{ kpc} < R < \frac{3.2}{1+z} h^{-1} \text{ Mpc}$, centered on the MCMF determined optical cluster center. We limit the use of DES Y3 weak-lensing data to clusters with a redshift below $z = 0.95$, leading to 688 cluster shear profiles extracted using a sample of 555,912 source galaxies. We show in Fig. 2 the constructed averaged matter profiles for confirmed MCMF SPT clusters. The discussion of relevant sources of uncertainty including cluster member contamination, miscentering of the shear profile, shear and photo- z calibration, halo mass modeling, and impact of large-scale structure can be found in the dedicated cluster cosmology methods paper for this combination of data [58].

C. High-redshift HST weak-lensing data

DES Y3 data loses its constraining power at higher redshift ($z \gtrsim 0.9$), and therefore, we include also HST weak-lensing data. Because of its high resolution imaging, it provides very precise targeted measurements of shear profiles for clusters at higher redshift. We use the previously analyzed weak-lensing measurements for 39 clusters in the redshift range $0.6 - 1.7$ in [52, 73–76]. For a further description of this component of the dataset, we refer the reader to these papers. The constructed averaged matter profiles are shown in Fig. 2.

IV. ABUNDANCE AND MASS CALIBRATION ANALYSES

In this section, we summarize the analysis method used for the abundance and mass calibration likelihoods. We introduce the observable–mass relations, describe the method used to obtain the mass information from weak-lensing measurements and then present the calibration of the weak-lensing mass to halo mass mean relation ($M_{\text{WL}}-M_{\text{halo}}$), which we use to model systematic uncertainties in the weak-lensing mass calibration. Finally, we describe the ingredients entering the pipeline, including the relevant likelihoods and priors on parameters. Our description follows the presentation in [58].

A. Observable Mass Relations

The observed tSZE significance $\hat{\zeta}$ is related to an unbiased or intrinsic detection significance ζ as [77]

$$P(\hat{\zeta}|\zeta) = \mathcal{N}\left(\sqrt{\zeta^2 + 3}, 1\right), \quad (14)$$

where $\mathcal{N}(a, \sigma)$ refers to a Gaussian distribution with mean a and standard deviation σ . The normal distribution arises from Gaussian noise in the survey maps. The bias correction factor of 3 accounts for noise bias introduced during the matched-filter search for peaks, considering three parameters: two for the location on the sky and one for the effective core radius of the tSZE signature.

The observable–mass relation, which links the mean intrinsic significance ζ to the halo mass M_{200c} used in describing the Halo Mass Function (HMF), is modeled as a power law in both mass and the redshift-dependent dimensionless Hubble parameter ($E(z) \equiv H(z)/H_0$)

$$\begin{aligned} \langle \ln \zeta \rangle = & \ln \zeta_0 + \zeta_M \ln \left(\frac{M_{200c}}{3 \cdot 10^{14} h^{-1} M_\odot} \right) \\ & + \zeta_z \ln \left(\frac{E(z)}{E(0.6)} \right). \end{aligned} \quad (15)$$

The parameters ζ_0 , ζ_M , and ζ_z represent the relation amplitude, mass trend, and redshift trend, respectively. The quantity M_{200c} denotes the halo mass within a region where the overdensity is 200 times the critical density at the cluster redshift. The pivot redshift and mass for the power law relation are set to $z = 0.6$ and $M_{200c} = 3 \cdot 10^{14} h^{-1} M_\odot$, respectively. Both observed and simulated clusters exhibit structural variations due to factors such as the time elapsed since their last major merger, which introduces scatter in the observable at any given mass and redshift. We model the intrinsic scatter of the intrinsic significance ζ around the mean relation as a log-normal distribution characterized by its mass and redshift independent RMS (root mean square) variation, $\sigma_{\ln \zeta}$.

As the various SPT surveys have different depths, which we characterize using the parameter γ_{field} , we rescale ζ_0 for each field with

$$\zeta_{0,\text{field}} = \gamma_{\text{field}} \zeta_0. \quad (16)$$

This variation in depth values affects the redshift evolution parameter ζ_z as well, therefore, we rescale it for each SPT survey. A more complete discussion of the survey information and the rescaling of parameters are listed in [58].

The noise in the observed richness $\hat{\lambda}$ of a particular cluster is represented by a Poisson sampling of an intrinsic richness λ . For $\lambda > 10$, it is valid to take a lognormal limit for the Poisson distribution, this leads to

$$P(\ln \hat{\lambda} | \ln \lambda) = \mathcal{N}(\ln \lambda, 1/\lambda). \quad (17)$$

As for detection significance, we use the same ansatz for modelling the intrinsic richness

$$\begin{aligned} \langle \ln \lambda \rangle = & \ln \lambda_0 + \lambda_M \ln \left(\frac{M_{200c}}{3 \cdot 10^{14} h^{-1} M_\odot} \right) \\ & + \lambda_z \ln \left(\frac{1+z}{1.6} \right). \end{aligned} \quad (18)$$

The richness–mass relation parameters λ_0 , λ_M , and λ_z correspond to the amplitude, mass trend, and redshift trend, respectively. The log-normal scatter of the intrinsic richness at a given mass and redshift around the mean richness is modeled using its RMS variation $\sigma_{\ln \lambda}$.

We utilize two distinct measurements of richness: DES Y3 data for clusters with redshift $z < 1.1$ and WISE data for clusters with redshift $z > 1.1$. Instead of aligning the two richness measurements within the overlapping redshift range, we establish separate observable–mass relations for each, and transition from DES-based to WISE-based richnesses at $z = 1.1$.

We calibrate the observable–mass relations in eqs. (15) and (18) with their intrinsic scatter empirically with weak-lensing data. A detailed description of the method is presented for DES Y3 elsewhere [58], but we briefly describe the main steps here.

The weak lensing (WL) observable for each cluster is the reduced tangential shear $g_t(R)$, which consists of a single shear profile that merges all tomographic bins of weak-lensing source galaxies. We use this observable to determine the weak-lensing halo mass (M_{WL}) by fitting the reduced shear profile to a Navarro-Frenk-White (NFW) profile [78] for the mass distribution within the radial range of $500 h^{-1} \text{ kpc} < R < \frac{3.2}{1+z} h^{-1} \text{ Mpc}$.

The inferred weak lensing mass M_{WL} is not an exact determination of the halo mass, because halos do not individually match perfectly to the NFW model. We account for differences between M_{WL} and the halo mass by establishing a $M_{WL} - M_{\text{halo}}$ mean relation [79–81]

$$\begin{aligned} \left\langle \ln \left(\frac{M_{WL}}{2 \cdot 10^{14} h^{-1} M_\odot} \right) \right\rangle = & \ln M_{WL_0} \\ & + M_{WL_M} \ln \left(\frac{M_{200c}}{2 \cdot 10^{14} h^{-1} M_\odot} \right), \end{aligned} \quad (19)$$

considering that M_{halo} is M_{200c} . $\ln M_{WL_0}$ is the logarithmic mass bias at $M_{200c} = 2 \cdot 10^{14} h^{-1} M_\odot$ and M_{WL_M} is the mass trend in this bias. The width of the lognormal scatter around the mean of M_{WL} is described as

$$\begin{aligned} \ln \sigma_{\ln WL}^2 = & \ln \sigma_{\ln WL_0}^2 \\ & + \sigma_{\ln WL_M}^2 \ln \left(\frac{M_{200c}}{2 \cdot 10^{14} h^{-1} M_\odot} \right). \end{aligned} \quad (20)$$

We generate synthetic cluster shear maps by applying the source redshift distribution, cluster miscentering, and cluster member contamination to halo mass maps from numerical simulations (see description below). These synthetic shear maps are then used to calibrate the free parameters of the $M_{WL} - M_{\text{halo}}$ relation [81].

Another way of parameterizing a parameter p with its standard deviation Δp is by splitting it as

$$p = \mathcal{N}(\bar{p}, \Delta p) = \bar{p} + \Delta p \mathcal{N}(0, 1), \quad (21)$$

where \bar{p} is the mean value and Δp is the uncertainty of the corresponding parameter p obtained from simulations. To accurately describe the redshift dependent

uncertainty of the logarithmic mass bias $\ln M_{\text{WL}_0}$, the scatter in this parameter, $\Delta \ln M_{\text{WL}_0}(z)$, is modeled as a linear combination of two redshift-dependent functions

$$\Delta \ln M_{\text{WL}_0}(z) = \Delta_1 \ln M_{\text{WL}_0}(z) + \Delta_2 \ln M_{\text{WL}_0}(z). \quad (22)$$

To establish a relationship between M_{WL} and the “gravity-only mass” M_{halo} we use mass maps from full-physics hydrodynamical simulations paired with gravity-only simulations with identical initial conditions. This approach includes the effect of baryons while leveraging the more robust halo mass function predictions from gravity-only simulations such as [82], including emulators like [83–85]. In practice, we apply this method to the Magneticum [86–89] and the Illustris TNG simulations [90–95]. The resulting constraints on the $M_{\text{WL}}-M_{\text{halo}}$ relation show some variation, which we attribute to uncertainties in modeling baryonic effects. To account for this, we inflate the uncertainty on all parameters of the $M_{\text{WL}}-M_{\text{halo}}$ relation by the additional uncertainty listed in Table 2 of [81], which, for the amplitude $\ln M_{\text{WL}_0}$, amounts to 0.02 or 2%. This level of systematic uncertainty is smaller than the current statistical uncertainty in the lensing dataset, ensuring that our analysis is not overly dependent on the accuracy of the Magneticum or Illustris TNG simulations in replicating the real universe.

HST-39 data is a targeted weak-lensing dataset for clusters. Given the selection of lensing sources via conservative color cuts, cluster member contamination was found to be negligible for this sample (e.g. [52]), which is why it is not modeled here. miscentring has been accounted for this sample as part of the mass bias estimation, as described in [52, 96], assuming isotropic miscentring (see [97] regarding the limitation of this assumption). The HST $M_{\text{WL}}-M_{\text{halo}}$ mean relation has only the first term in eq. (19) ($\ln M_{\text{WL}_0}$) and its scatter ($\ln \sigma_{\ln \text{WL}_0}^2$) [73].

In summary, the observable–mass relations (ζ –mass- z and λ –mass- z) along with the $M_{\text{WL}}-M_{\text{halo}}$ relation are power law relations as described in Eqs. (15), (18), and (19), respectively. All have an intrinsic scatter which we model as log-normal. We establish a covariance matrix between all of the scatter parameters, and account for possible correlations among all pairs. We introduce the parameters $\rho_{\zeta, \text{WL}}$, $\rho_{\zeta, \lambda}$, and $\rho_{\text{WL}, \lambda}$ for significance–weak-lensing, significance–richness, and weak-lensing–richness correlations, respectively.

B. Cluster Abundance Likelihood

We use Bayes’ theorem to infer values of cosmological parameters \mathbf{p} assuming a cluster population model. We model the likelihood of cluster population as a Poisson realization of the halo observable function (HOF). The

log-likelihood is given by

$$\begin{aligned} \ln \mathcal{L}(\mathbf{p}) = & \sum_i \ln \left. \frac{d^4 N(\mathbf{p})}{d\hat{\zeta} d\hat{\lambda} d\mathbf{g}_t dz} \right|_{\hat{\zeta}_i, \hat{\lambda}_i, \mathbf{g}_t, i, z_i} \\ & - \int \dots \int d\hat{\zeta} d\hat{\lambda} d\mathbf{g}_t dz \frac{d^4 N(\mathbf{p})}{d\hat{\zeta} d\hat{\lambda} d\mathbf{g}_t dz} \Theta_s(\hat{\zeta}, \hat{\lambda}, z) + \text{const.} \end{aligned} \quad (23)$$

Here Θ_s is the survey selection function and it is defined in terms of the lower limit thresholds we impose in the observables $\hat{\zeta}$, $\hat{\lambda}$, and z as described in Sec. III A. The lensing data are tangential shear profiles \mathbf{g}_t . The differential cluster abundance in observable space is defined as

$$\begin{aligned} \frac{d^4 N(\mathbf{p})}{d\hat{\zeta} d\hat{\lambda} d\mathbf{g}_t dz} = & \int d\Omega_s \iiint dM d\zeta d\lambda dM_{\text{WL}} \\ & P(\mathbf{g}_t | M_{\text{WL}}, \mathbf{p}) P(\hat{\zeta} | \zeta) P(\hat{\lambda} | \lambda) \\ & P(\zeta, \lambda, M_{\text{WL}} | M, z, \mathbf{p}) \\ & \frac{d^2 N(\mathbf{p}, M, z)}{dM dV} \frac{d^2 V(\mathbf{p}, z)}{dz d\Omega_s}, \end{aligned} \quad (24)$$

where $\frac{d^2 N(\mathbf{p}, M, z)}{dM dV}$ is the halo mass function, $\frac{d^2 V(\mathbf{p}, z)}{dz d\Omega_s}$ is the differential volume, and Ω_s is the survey footprint. The relations between observed and intrinsic parameters, $P(\hat{\zeta} | \zeta)$ and $P(\hat{\lambda} | \lambda)$, are defined in Eqs. (14) and (17). $P(\zeta, \lambda, M_{\text{WL}} | M, z, \mathbf{p})$ follows from Eqs. (15), (18), (19) and (20). Finally, the lensing likelihood $P(\mathbf{g}_t | M_{\text{WL}}, \mathbf{p})$ is defined as a product of independent Gaussian probabilities in each radial bin i of the tangential reduced shear profile of a given cluster

$$P(\mathbf{g}_t | M_{\text{WL}}, \mathbf{p}) = \prod_i \left(\sqrt{2\pi} \Delta g_{t,i} \right)^{-1} e^{-\frac{1}{2} \left(\frac{g_{t,i} - g_{t,i}(M_{\text{WL}}, \mathbf{p})}{\Delta g_{t,i}} \right)^2}, \quad (25)$$

with $\Delta g_{t,i}$ as the shape noise. Again, more discussion can be found in the cosmological methods paper [58].

C. Pipeline

The likelihood is implemented as a python module in the framework of **CosmoSIS**¹ [98]. We explore the parameter space using the nested sampler **MultiNest** [99–101]. For the purpose of comparison we include an analysis with CMB anisotropies from Planck 2018 [7] and BAO data from BOSS DR12 [8]. Note that an analysis of IDM–DR with Planck and BAO data was previously performed in [41, 43]. In this work, we consider the following dataset combinations

¹ <https://cosmosis.readthedocs.io/>

TABLE I. Priors on fitting parameters in the analysis of the abundance of SPT clusters with DES Y3 and HST weak-lensing data. The lensing model parameters are informed by priors derived from simulations, as detailed in Sec. IV A, and inform the empirical calibration of the other observable–mass relations. \mathcal{U} represents a flat distribution, and $\mathcal{N}(a, \sigma)$ refers to a Gaussian distribution with mean a and standard deviation σ .

Parameter	Description	Informative Prior
DES Y3 cluster lensing		
$\Delta_1 \ln M_{\text{WL}_0}(z)$	bias uncertainty as $f(z)$	$\mathcal{N}(0, 1)$
$\Delta_2 \ln M_{\text{WL}_0}(z)$	bias uncertainty as $f(z)$	$\mathcal{N}(0, 1)$
M_{WL_M}	mass trend of bias	$\mathcal{N}(1.029, 0.006)$
$\ln \sigma_{\text{in WL}_0}^2$	normalization of scatter	$\mathcal{N}(0, 1)$
$\sigma_{\text{in WL}_M}^2$	mass trend of scatter	$\mathcal{N}(-0.226, 0.040)$
HST cluster lensing		
$\ln M_{\text{WL}_0}$	amplitude of bias	$\mathcal{N}(0, 1)$
$\ln \sigma_{\text{in WL}_0}^2$	amplitude of scatter	$\mathcal{N}(0, 1)$
tSZE ζ -mass- z parameters		
$\ln \zeta_0$	amplitude	$\mathcal{U}(0.3, 2.0)$
ζ_M	mass trend	$\mathcal{U}(1.2, 2.0)$
ζ_z	redshift trend	$\mathcal{U}(-1.0, 1.5)$
$\sigma_{\text{in } \zeta}$	intrinsic scatter	$\mathcal{U}(0.05, 0.5)$
γ_{ECS}	depth of SPTpol ECS	$\mathcal{U}(0.9, 1.2)$
DES λ -mass- z parameters (used for $z < 1.1$)		
$\ln \lambda_0$	amplitude	$\mathcal{U}(3.0, 4.0)$
λ_M	mass trend	$\mathcal{U}(0.7, 1.5)$
λ_z	redshift trend	$\mathcal{U}(-1.0, 0.8)$
$\sigma_{\text{in } \lambda}$	intrinsic scatter	$\mathcal{U}(0.05, 0.5)$
WISE λ -mass- z parameters (used for $z > 1.1$)		
$\ln \lambda_0$	amplitude	$\mathcal{U}(3.0, 5.0)$
λ_M	mass trend	$\mathcal{U}(0.7, 1.5)$
λ_z	redshift trend	$\mathcal{U}(-4.0, 0.8)$
$\sigma_{\text{in } \lambda}$	intrinsic scatter	$\mathcal{U}(0.05, 0.5)$
Correlation coefficients		
$\rho_{\zeta, \text{WL}}$	tSZE ζ –WL	$\mathcal{U}(-0.5, 0.5)$
$\rho_{\zeta, \lambda}$	tSZE ζ – λ	$\mathcal{U}(-0.5, 0.5)$
$\rho_{\text{WL}, \lambda}$	WL– λ	$\mathcal{U}(-0.5, 0.5)$
Cosmology		
ξ_{DR}	DR temperature ratio	$\mathcal{U}(0.001, 0.5)$
Ω_m	matter density parameter	$\mathcal{U}(0.1, 0.5)$
$\Omega_b h^2$	baryon density parameter	$\mathcal{N}(0.02236, 0.00015)$
h	Hubble parameter	$\mathcal{N}(0.7, 0.05)$
$\ln 10^{10} A_s$	amplitude of $P(k)$	$\mathcal{U}(2.0, 5.0)$
n_s	scalar spectral index	$\mathcal{N}(0.9649, 0.0044)$

- SPT-cluster \times WL: a sample of 1,005 tSZE selected clusters from SPT surveys [53–56] confirmed with the MCMF algorithm and complemented with weak-lensing mass information from DES Y3 [49–51] and HST [52] as described in Sec. III
- CMB+BAO: CMB data from Planck 2018 (TT, TE, EE + lowE) [7] and BAO measurements from BOSS DR12 [8]

For SPT-clusters \times WL we employ the full log-likelihood containing both the cluster abundance and the weak lensing mass calibration simultaneously as given in Eq. (23). We list all priors for fit parameters in Ta-

ble I. As discussed in Sec. IV A, we employ informative, Gaussian priors on the weak-lensing mass calibration parameters introduced in Eqs. (19)–(20), for DES Y3 data and HST data separately as described in Sec. IV A. The observable–mass parameters introduced in Eqs. (14)–(18) are fully constrained by the mass calibration likelihood. We apply flat and uninformative priors on them. As explained in Sec. IV A, we consider two richness–mass relations for the datasets (DES Y3 and WISE) at two different redshift regimes. Therefore, we apply priors for them separately. We then define flat priors for the correlation coefficients in the covariance matrix as mentioned in the previous section with a range which ensures that the resulting correlation matrix remains non-singular across all parameter combinations. Finally, we apply priors for cosmological parameters. When performing an analysis for galaxy cluster abundance only, we define Gaussian priors on $\Omega_b h^2$ and n_s from Planck results and a wide prior on h covering all current available constraints [$\mathcal{N}(0.7, 0.05)$]. This is because clusters are not sensitive to the density of baryons and the spectral index, and they do not provide sufficient constraining power on the Hubble parameter by themselves. When including CMB and BAO data in the analysis, we apply only flat priors on the previously mentioned parameters, and marginalize over the optical depth of reionization τ as well.

The upper bound in the flat prior for ξ_{DR} arises from Planck constraints on ΔN_{eff} for non-interacting DR, translated to ξ_{DR} via Eq. (4) (see also [41, 43]). We note that the precise choice is irrelevant for our results, because the constraints on DR interacting with DM that we derive in this work are far below this upper limit. Moreover, following [44], we focus on the tight coupling limit, where the interaction intensity is set to a sufficiently high value that our results are independent of its exact value. Specifically, we use $\log_{10}[a_{\text{dark}}/\text{Mpc}^{-1}] = 8$. We also fix the IDM fraction to $f_{\text{IDM}} = 10\%$ to simplify the analysis and avoid parameter projection effects. Note that the matter power spectrum and HMF are relatively insensitive to this precise value within the relevant mass regime of the cluster sample we study here; more discussion on this choice can be found in [44]. A complementary analysis including both a_{dark} and f_{IDM} as free parameters is provided in Appendix A.

V. RESULTS

We present constraints on cosmological parameters within the IDM–DR model from galaxy cluster abundance based on the SPT sample with weak-lensing informed masses from DES Y3 and HST as described in Sec. III. In addition, we discuss the interplay of the IDM–DR model with massive neutrinos and provide an assessment of the scale-dependent sensitivity of cluster abundance measurements in comparison to other datasets. Our full results including also all

observable–mass parameters entering the analysis are shown in Appendix B.

A. Cosmological constraints in the IDM–DR model

We show our main results in Fig. 3² and summarize constraints on model parameters in Table II. On the left side of Fig. 3, we present marginalized two-dimensional posteriors for σ_8 , Ω_m and ξ_{DR} . In IDM–DR, SPT-clusters×WL prefer slightly decreasing values of σ_8 when increasing ξ_{DR} . A similar tendency was also pointed out when considering CMB and galaxy clustering data in [41, 43]. This can be associated to a suppression of the fluctuation amplitude due to the IDM–DR interaction, that shifts to larger scales with increasing ξ_{DR} . From SPT-clusters×WL we obtain an upper bound on the DR temperature ratio $\xi_{\text{DR}} < 0.166$ at 95% credibility. Adding clusters to CMB and BAO measurements yields a significant improvement in the joint constrain on the temperature ratio with $\xi_{\text{DR}} < 0.098$ (95% credibility), compared to CMB+BAO alone $\xi_{\text{DR}} < 0.130$ (95% credibility).

The upper bound on ξ_{DR} translates into an upper bound on ΔN_{eff} , as defined in Eq. (4). In an $SU(3)$ theory, this leads to $\Delta N_{\text{eff}} < 0.003$, which is about three times better than the constraint from CMB+BAO alone. Note that this constraint applies only to DR tightly coupled to IDM, and in the fluid limit. For reference, we note that this bound is ~ 100 times stronger than the one on free-streaming and non-interacting extra radiation from Planck CMB data combined with BAO ($\Delta N_{\text{eff}} < 0.28$ [18]). This can be explained by the strong impact of the DR density on the HMF due to its tight coupling to IDM and the associated effect on the matter power spectrum, which is absent for non-interacting and free-streaming DR.

As can be seen in Table II, constraints on the parameters Ω_m and σ_8 from SPT-clusters×WL within IDM–DR are broadly consistent with those obtained assuming ΛCDM [5]. On the right side of Fig. 3, we show the constraints on S_8 and S_8^{opt} . In the light of the S_8 tension, we note that our results are consistent with both early time measurements (e.g. from Planck [7]) and late time measurements (e.g. cosmic shear measurements in the joint analysis of KiDS and DES data [13]) within the uncertainty range. Notably, when combining cluster data with CMB and BAO measurements, the constraining power on S_8 is two times improved compared to CMB+BAO only. Note that cluster abundance measurements are most sensitive to the parameter combination $S_8^{\text{opt}} = \sigma_8(\Omega_m/0.3)^{0.2}$. As shown in the lower right

part of Fig. 3, the degeneracy between S_8^{opt} and Ω_m is completely broken. From Table II, we see that the uncertainty on S_8^{opt} is considerably reduced by about 40% compared to S_8 when considering SPT-clusters×WL within IDM–DR. Moreover, the error on both S_8^{opt} as well as S_8 is only degraded by about 10% within IDM–DR as compared to ΛCDM . We further highlight the finding that the posteriors for σ_8 , S_8^{opt} and S_8 feature a tail towards lower values, see Fig. 3, leaving room for addressing the S_8 tension within IDM–DR.

Interestingly, we also find a slight preference for a non-zero value of ξ_{DR} . While not being statistically significant, we checked that this hint is robust to several variations in the analysis, such as including neutrino masses as in Sec. VB. We also note that it persists when including additionally a_{dark} and f_{IDM} as free model parameters, see Appendix A. This is intriguing since both of these parameters become irrelevant in the limit $\xi_{\text{DR}} \rightarrow 0$, leading to an increased volume of parameter space entering the marginalization, and therefore potentially a preference of the posterior for smaller values of ξ_{DR} . Nevertheless, the hint for non-zero ξ_{DR} is also present in this case. It will be interesting to follow up on this hint with future cluster abundance measurements and complementary datasets.

B. Effect of neutrinos

Massive neutrinos suppress structure formation below their free-streaming scale. As previously discussed in [44], this suppression is distinct from the one within the IDM–DR model, occurring at a different length scale. Additionally, neutrino masses also affect the background evolution after recombination, in contrast to IDM–DR. Although the effect of massive neutrinos is distinct from the one of the IDM–DR model, we carry out an analysis when adding the sum of neutrino masses as a free parameter, in order to check whether including it affects the constraint on ξ_{DR} .

We present a comparison between the case of fixing the sum of neutrino masses to $\sum m_\nu = 0.06 \text{ eV}$ ³ and the case of letting it free in Fig. 4. We find that the sum of neutrino masses does not affect constraints on ξ_{DR} significantly. Importantly, there is no degeneracy between $\sum m_\nu$ and ξ_{DR} . This agrees with the expectation from the sensitivity forecast of cluster abundance measurements from CMB-S4 and SPT-3G in [44]. In turn, we also find that the neutrino mass bounds are not significantly degraded within IDM–DR as compared to ΛCDM . We obtain an upper bound $\sum m_\nu < 0.096 \text{ eV}$ when considering the data combination SPT-clusters×WL+CMB+BAO. This upper limit

² All plots are generated using GetDist <https://getdist.readthedocs.io/en/latest/>.

³ This is the minimum value for neutrino masses in the normal hierarchy as indicated by neutrino flavour oscillation experiments, for a review see [102].

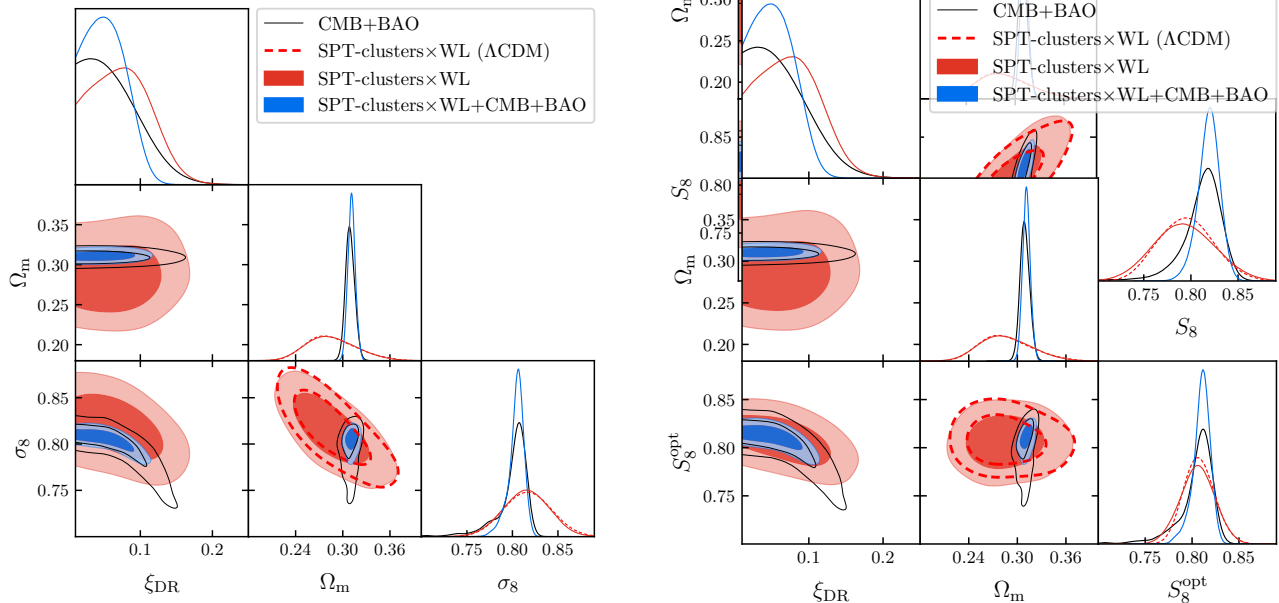


FIG. 3. Posteriors obtained from analysing the IDM–DR model with galaxy cluster abundance measurements (SPT tSZE selected and MCMF optical/NIR confirmed clusters from the SPT-SZ, SPTpol ECS, and SPTpol 500d surveys) combined with weak lensing data (DES Y3 and HST) for mass calibration (red). We show the results also in combination with CMB (Planck 2018) and BAO (BOSS DR12) measurements (blue), and compare with results for CMB+BAO alone (black lines). We include the analysis of cluster data within the Λ CDM model from [5] (dashed red lines) for comparison. On the left, constraints on ξ_{DR} , Ω_m and σ_8 at 68% credible intervals (CI) and 95% upper limits are shown. On the right, we display constraints for the commonly used parameter combination $S_8 = \sigma_8(\Omega_m/0.3)^{0.5}$, and $S_8^{\text{opt}} = \sigma_8(\Omega_m/0.3)^{0.2}$ because it is the combination galaxy cluster counts are most sensitive to.

TABLE II. Parameter constraints for the IDM–DR and Λ CDM models (marginal mean at 68% CI, or upper limit at 95% CI). $S_8^{\text{opt}} = \sigma_8(\Omega_m/0.3)^{0.2}$ corresponds to the combination that cluster counts are most sensitive to. Results for Λ CDM are from [5].

Dataset	ξ_{DR}	ΔN_{eff}^a	Ω_m	σ_8	S_8	S_8^{opt}
IDM–DR model						
CMB+BAO	< 0.130	< 0.010	0.312 ± 0.005	0.801 ± 0.019	0.817 ± 0.021	0.809 ± 0.020
SPT-clusters×WL	< 0.166	< 0.027	0.285 ± 0.032	0.815 ± 0.025	0.793 ± 0.032	0.805 ± 0.018
SPT-clusters×WL+CMB+BAO	< 0.098	< 0.003	0.311 ± 0.005	0.804 ± 0.009	0.819 ± 0.011	0.810 ± 0.009
Λ CDM model						
SPT-clusters×WL	0.286 ± 0.032	0.817 ± 0.026	0.795 ± 0.029	0.807 ± 0.016

^a Note that here ΔN_{eff} refers to DR that is tightly coupled to IDM and acts as a fluid. This explains the difference to the usually considered case of free-streaming, non-interacting extra radiation.

is more stringent compared to $\sum m_\nu < 0.13 \text{ eV}$ obtained from CMB+BAO [7].

C. Assessing the scale sensitivity and comparison to other probes

When testing for potential deviations from the Λ CDM paradigm it is instructive to quantify to which range of length-scales a given measurement is dominantly sensitive to, and to which level of precision. Here we address this question for cluster number counts, CMB+BAO

as well as weak-lensing shear data and galaxy clustering. In Fig. 5 we show an approximate assessment of the scale sensitivity of SPT-clusters×WL, CMB+BAO as well as their combination. We determine the linear matter power spectrum for each set of model parameters contained in the converged samples within our analysis based on the IDM–DR framework. We then display the variation of the power spectrum around the mean value within 1σ for each wave number, with $\Delta P(k) = (P(k)_{\text{mean}} \pm \text{stdev})/P(k)_{\text{mean}}$.

As expected, we find that cluster abundance measurements are most sensitive to larger scales, approximately

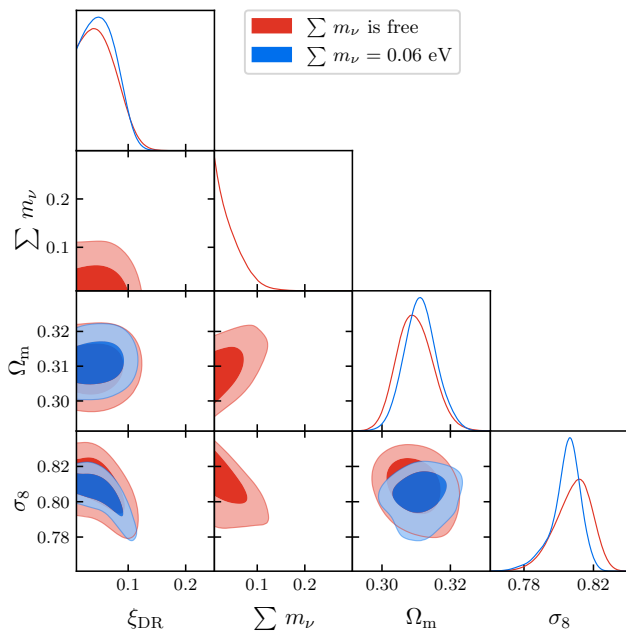


FIG. 4. Comparison between the combined posteriors obtained when fixing the sum of neutrino masses to $\sum m_\nu = 0.06$ eV (blue), and when marginalizing over it (red). We show results for the combination SPT-clusters \times WL+CMB+BAO.

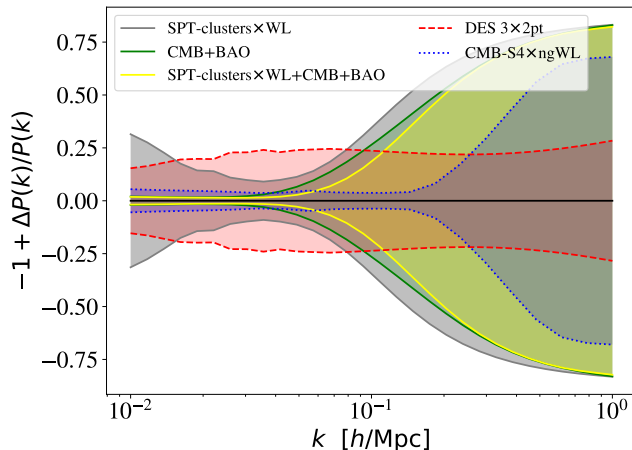


FIG. 5. Variance (1σ region) of the matter power spectrum as a function of wave number k for different cosmological probes, inferred from converged samples of model parameters. We set $\Delta P(k) = P(k) \pm \text{stdev}$, where $P(k)$ is the mean value of the matter power spectrum. Apart from SPT-clusters \times WL, CMB+BAO and their combination, we also show corresponding results from weak lensing shear measurements (DES 3×2 pt results computed in the Λ CDM model from [12]) as well as a forecast for future cluster abundance measurements combined with next-generation weak-lensing data (CMB-S4 \times ngWL, see [44] for details).

$k \in [0.02, 0.08] h/\text{Mpc}$. This means they are less affected by the non-linear evolution of the universe and small scale effects like baryonic feedback. CMB+BAO data are highly sensitive to very large scales with small uncertainty. We note that the sensitivity of CMB+BAO measurements improves when combining them with galaxy cluster data (as shown in yellow in Fig. 5) in particular on scales $k \sim 0.1 h/\text{Mpc}$ that are particularly relevant for probing IDM–DR. All of these datasets loose sensitivity when going to smaller scales, in this case $k \gtrsim 0.1 h/\text{Mpc}$. This explains why clusters and CMB+BAO are able to constrain the value of ξ_{DR} , which, as discussed and shown in Fig. 1 in [44], is responsible for defining the scale at which the suppression in the matter power spectrum sets in. The higher the value of ξ_{DR} , the more the suppression shifts towards larger scales. This scale-dependence also explains the relatively weak sensitivity of both clusters as well as CMB+BAO data to the IDM–DR model parameters a_{dark} and f_{IDM} . Constraining them would require including datasets that are more sensitive to smaller scales, such as cosmic shear measurements, galaxy clustering, and their correlation (see DES Y3 3×2 pt results within Λ CDM in Fig. 5). A combined analysis of cluster abundance with cosmic shear is thus a promising avenue for covering the entire IDM–DR model parameter space, with clusters providing sensitivity to the DR temperature ξ_{DR} and 3×2 pt functions to the fraction of interacting DM f_{IDM} and/or the interaction strength a_{dark} . We also note the sensitivity of cluster data will improve significantly with ongoing and future surveys. This is shown in blue in Fig. 5 for the case of clusters that could be detected with a CMB-S4 survey combined with next-generation weak-lensing data from Euclid or Rubin. This is owed to the larger size of the sample which will have more than an order of magnitude compared to current samples, and to the improved systematical uncertainties, more discussion can be found in [44].

VI. CONCLUSION

In this work, we derive constraints on interactions of dark matter with dark radiation using the tSZE detected SPT cluster sample consisting of 1,005 clusters, with weak lensing measurements for 688 clusters from DES Y3 and for 39 clusters from HST, and employing the galaxy cluster abundance analysis framework developed in [5, 58]. The cosmological IDM–DR model we consider has been discussed as a potential solution to the S_8 tension, and is part of the ETHOS framework that can describe a wide generic class of underlying particle physics models. A prominent example is a dark sector featuring a weakly coupled, unbroken non-Abelian $SU(N)$ gauge symmetry. The model is characterized by three parameters, being the temperature ratio ξ_{DR} , the fraction of interacting dark matter f_{IDM} , and the IDM–DR interaction strength a_{dark} . The latter is related to the dark fine structure constant within the microscopic realization pro-

vided by the $SU(N)$ model. Galaxy cluster abundance data are mainly sensitive to ξ_{DR} .

We find that the tSZE cluster sample obtained by SPT complemented with mass information from DES Y3 and HST yields an upper bound on $\xi_{\text{DR}} < 17\%$ (95% credibility). When combining these cluster abundance data with CMB data from Planck 2018 and BAO measurements from BOSS DR12, we obtain the most stringent to date upper bound $\xi_{\text{DR}} < 10\%$ (95% credibility). This bound is 30% tighter than the one from CMB+BAO alone. Note that this translates into an improvement in the upper bound on the DR energy density by a factor of three when adding clusters to CMB+BAO data.

Within the IDM–DR model, the cluster data set considered in this work yields $S_8 = 0.793 \pm 0.032$ (68% CI) as well as $S_8^{\text{opt}} = \sigma_8(\Omega_m/0.3)^{0.2} = 0.805 \pm 0.018$ for the combination clusters are most sensitive to. This is consistent both with Planck CMB as well as with weak lensing shear measurements. Furthermore, slightly lower values of S_8 can be accommodated within the IDM–DR model compared to Λ CDM, leaving room for addressing the S_8 tension within this class of models. Interestingly, we find a slight preference for a non-zero value of ξ_{DR} , that is however statistically not significant. Yet, it occurs both when restricting the model to the limit for which DR and IDM are tightly coupled, as well as in the general case.

Future galaxy cluster abundance data such as from CMB-S4 combined with mass information from Euclid or Rubin weak-lensing measurements will significantly increase the sensitivity to ξ_{DR} . Furthermore, a combined analysis with weak lensing shear data could provide enhanced sensitivity to the model parameters f_{IDM} and a_{dark} , and would be instrumental for scrutinizing the question of whether IDM–DR models can address the potential S_8 tension.

In summary, our results show that weak-lensing informed galaxy cluster abundance measurements are sensitive to fundamental properties of DM, testing its cold and collisionless nature with unprecedented sensitivity, and in a way that is complementary to both CMB and cosmic shear measurements.

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APPENDIX A: Varying a_{dark} and f_{IDM}

The main results presented in this work are based on considering the tight coupling limit (meaning $\log_{10}[a_{\text{dark}}/\text{Mpc}^{-1}] > 8$) and fixing the value of $f_{\text{IDM}} = 0.1$. This choice is motivated by the observation that the halo mass function (HMF) is mostly sensitive to the value of the DR temperature within the galaxy cluster mass range probed by SPT, as explained in [44]. Nevertheless, for completeness, we also perform an analysis where a_{dark} and f_{IDM} are kept as free parameters. We include a flat prior $\mathcal{U}(0.0, 10.0)$ for $\log_{10}[a_{\text{dark}}/\text{Mpc}^{-1}]$, and $\mathcal{U}(0.001, 1.0)$ for f_{IDM} .

As shown in Fig. 6, cluster abundance data as well as CMB+BAO data are only very weakly sensitive to a_{dark}

and f_{IDM} . However, importantly, we find that the constraint on the DR temperature ξ_{DR} is consistent with the results obtained when fixing the other two parameters as presented in Sec. V A. This implies that the constraints on ξ_{DR} are robust with respect to variations in the ETHOS parameters a_{dark} and f_{IDM} over a wide range. Moreover, interestingly, the slight preference for non-zero ξ_{DR} observed in the main text is also present when allowing a_{dark} and f_{IDM} to vary.

We note that we additionally performed an analysis for which we fixed $\log_{10}[a_{\text{dark}}/\text{Mpc}^{-1}] = 8$, i.e. within the tight coupling limit, but include f_{IDM} as an additional free parameter, finding very similar results to those shown in Fig. 6.

APPENDIX B: Full triangle plot

In Fig. 7 we show an extended version of Fig. 3 for the constraints obtained from SPT-clusters×WL data within the IDM–DR model. Fig. 7 includes all cosmological model parameters as well as the full set of parameters with flat-priors entering the cluster abundance and weak-lensing mass-calibration likelihoods that are jointly varied in our main analysis setup. The plot shows also consistent results with ΛCDM (contours from [5]).

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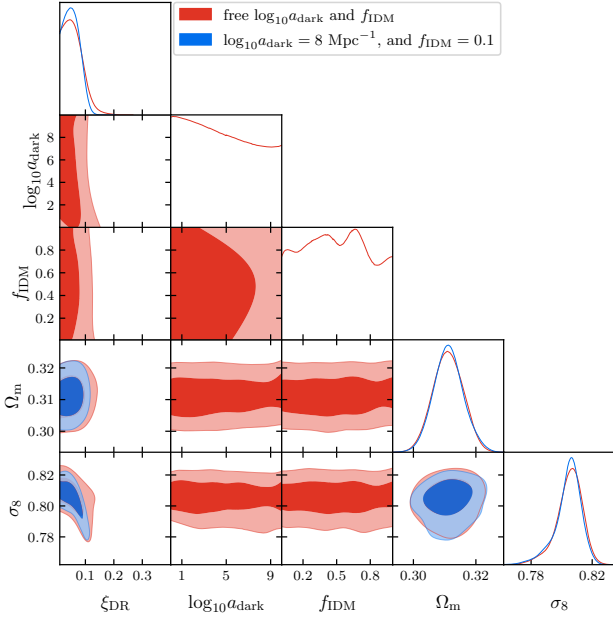


FIG. 6. Posteriors of IDM–DR model parameters obtained from SPT-clusters×WL+CMB+BAO data when allowing for variations in a_{dark} and f_{IDM} (red), compared to the case discussed in the main text where both parameters are fixed (blue).

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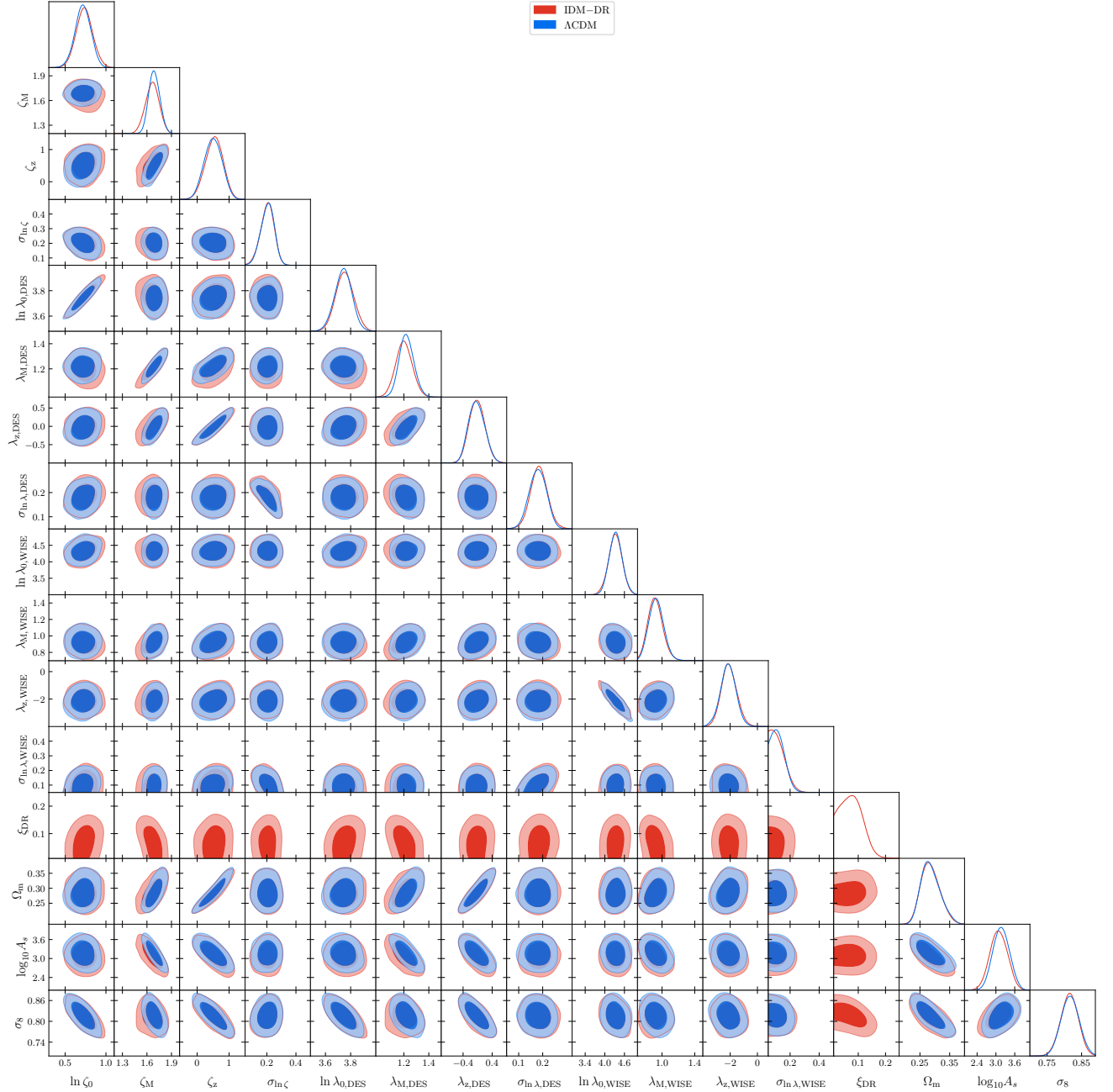


FIG. 7. Posterior distribution of all parameters with a flat prior employed in the analysis of SPT-clusters \times WL within IDM-DR model (red) and Λ CDM model (blue) from [5].

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