

A high-Q microwave dielectric resonator for axion dark matter haloscopes

R. Di Vora,^{1,2, a)} D. Alesini,³ C. Braggio,^{4,1, b)} G. Carugno,¹ N. Crescini,^{5, c)} D. D'Agostino,⁶ D. Di Gioacchino,³ P. Falferi,^{7,8} U. Gambardella,⁶ C. Gatti,³ G. Iannone,^{9,6} C. Ligi,³ A. Lombardi,⁵ G. Maccarrone,³ A. Ortolan,⁵ R. Pengo,⁵ A. Rettaroli,^{3,10} G. Ruoso,⁵ L. Taffarello,¹ and S. Tocci³

¹⁾*INFN, Sezione di Padova, Padova, Italy*

²⁾*Dipartimento di Scienze Fisiche, della Terra e dell'Ambiente, Università di Siena, via Roma 56, 53100 Siena, Italy*

³⁾*INFN, Laboratori Nazionali di Frascati, Frascati, Roma, Italy*

⁴⁾*Dipartimento di Fisica e Astronomia, Padova, Italy*

⁵⁾*INFN, Laboratori Nazionali di Legnaro, Legnaro, Padova, Italy*

⁶⁾*INFN, Sezione di Napoli, Napoli, Italy*

⁷⁾*Istituto di Fotonica e Nanotecnologie, CNR Fondazione Bruno Kessler, I-38123 Povo, Trento, Italy*

⁸⁾*INFN, TIFPA, Povo, Trento, Italy*

⁹⁾*Dipartimento di Fisica E.R. Caianiello, Fisciano, Salerno, Italy*

¹⁰⁾*Dipartimento di Matematica e Fisica, Università di Roma Tre, Roma, Italy*

(Dated: 13 January 2022)

The frequency band 1-15 GHz provides exciting prospects for resonant axion haloscopes as indicated by cosmological and astrophysical arguments. Among the challenges currently addressed to reach the required sensitivity, the development of high quality factor cavities that tolerate multi-Tesla fields plays a central role. We report a 3D resonator based on a right circular copper cavity with hollow cylinders that confine higher order modes around the cylinder axis. Its effective volume at 10.3 GHz is $3.4 \cdot 10^{-2}$ liters, and under an 8 T-field we measured an internal quality factor of more than 9 millions. These parameters demonstrate the potential of this unique resonator to probe galactic dark matter axion at remarkable scan rates of 15 MHz/day when the cavity is readout by a quantum-limited receiver.

I. INTRODUCTION

According to cosmological models, galaxies are surrounded by dark matter (DM) halos that extend far beyond the visible structures. Though DM halos cannot be observed directly, their existence is inferred through observations of their gravitational effects, and play a central role in our understanding of the Universe evolution¹. The nature of particle dark matter is unknown, and this is one of the most compelling problems in fundamental physics.

Unrelated to this open problem, a hypothetical particle named axion was introduced in the late 80s to answer a fundamental question in particle physics, namely why in the strong sector the charge-parity symmetry is not violated^{2,3}. However, physicists later realized that sufficiently light axions could also solve the DM problem⁴. As for detectability, their dominant interaction is via two-photon coupling in the Lagrangian

$$\mathcal{L} = g_{a\gamma\gamma} a \mathbf{E} \cdot \mathbf{B}, \quad (1)$$

where $g_{a\gamma\gamma}$ is the axion-photon coupling constant, proportional to the axion mass m_a , a is the axion field, \mathbf{E} and \mathbf{B} are the local electric and magnetic field vectors, respectively. The related process lifetime exceeds the age of the Universe, but P. Sikivie early in the 80's showed that if the field \mathbf{B} in eq. I was provided by a laboratory magnet, it would have been possible

to probe the existence of this dark matter candidate using microwave cavities permeated by multi-Tesla fields⁵. This is the basic principle of the so-called haloscope, that detects axions through their conversion to photons in a microwave cavity using the inverse Primakoff effect and a large magnetic field^{6,7}.

As implicit in eq. I, the axion to photon conversion probability in a cavity is proportional to $\int_V \mathbf{E}_{mnl} \cdot \mathbf{B}_0 dV$, where \mathbf{E}_{mnl} is the electric field of the resonant mode, \mathbf{B}_0 is the applied magnetic field and V is the cavity volume. Thus if the cavity frequency ω_c matches the axion frequency $\omega_a = m_a c^2 / \hbar$, the signal power in the cavity is given by

$$P_a = g_{a\gamma\gamma}^2 \frac{\rho_a}{m_a^2} \omega_c B_0^2 C_{mnl} V \frac{Q_c Q_a}{Q_c + Q_a}, \quad (2)$$

where C_{mnl} is the form factor that accounts for the spatial overlap between the external magnetic field and a given cavity mode. If Q_0 is the unloaded quality factor of the mode, Q_c in eq. 2 is given by $Q_c = Q_0 / (1 + \beta)$, with β coupling coefficient of the receiver antenna to the axion-sensitive cavity mode.

As the required intense magnetic field values are available in solenoidal configurations, the right circular cylindrical cavity is typically the chosen geometry to maximize C_{mnl}

$$C_{mnl} = \frac{|\int d^3x \mathbf{B} \cdot \mathbf{E}_{mnl}(\mathbf{x})|^2}{B_0^2 V \int d^3x \epsilon_r(\mathbf{x}) |\mathbf{E}_{mnl}(\mathbf{x})|^2},$$

where ϵ_r is the relative dielectric constant inside the cavity volume. Pertinent to the cavity resonator is also the last ratio in eq. 2, which expresses the dependence of the conversion power from the reduced quality factor of the system⁸ $Q_a^{-1} + Q_c^{-1}$, with $Q_a^{-1} \simeq 10^{-6}$ axion signal relative linewidth⁷. When no signature of the axion is observed, exclusion limits

^{a)}Electronic mail: divora@pd.infn.it

^{b)}Electronic mail: caterina.braggio@unipd.it

^{c)}Present address: University Grenoble Alpes, CNRS, Grenoble INP, Institut Néel, 38000 Grenoble, France.

are reported at a certain confidence level, assuming that axions saturate the dark matter density $\rho_a = \rho_{\text{DM}}$, whose value is typically taken as $\rho_{\text{DM}} = 0.45 \text{ GeV/cm}^3$ [7]. For a typical haloscope experiment the power given by eq. 2 is in the range from 10^{-23} to 10^{-25} W, several orders of magnitude below the electronic noise power P_N in state-of-the-art cryogenic receivers. For instance, the QUAX- $a\gamma$ haloscope⁹ recently set a 90% confidence level limit to the axion-photon coupling $g_{a\gamma\gamma} < 0.766 \cdot 10^{-13} \text{ GeV}^{-1}$ at $m_a \simeq 43 \mu\text{eV}$ ($\omega_c = 2\pi 10.4 \text{ GHz}$), using a cavity with $V = 80.56 \text{ cm}^3$, $C_{mnl} = 0.69$, unloaded quality factor $Q_0 = 72000$, coupling coefficient $\beta = 1$. Altogether, the expected power for these parameters at the benchmark for KSVZ models^{10,11} is $P_a = 3.8 \times 10^{-24}$ W. Under the assumption that the noise power follows a Gaussian distribution, as is the case for Johnson-Nyquist noise, experiments rely on the radiometer equation¹² $\Sigma = P_a \sqrt{t \Delta\nu} / P_N$ to set the integration time t to satisfy $\Sigma \simeq 2$ over a bandwidth $\Delta\nu$.

The Axion Dark Matter eXperiment (ADMX), based on a 136 l-volume cylindrical cavity, has excluded axions with masses $2.7 \mu\text{eV} - 4.2 \mu\text{eV}$ (650 MHz–1025 MHz)^{13,14} down to benchmark DFSZ axion models¹⁵. However, existing astrophysical and cosmological constraints within axion parameter space motivate the construction of new detectors to search for heavier axions⁷, in the two decade window $10 \mu\text{eV} - 1 \text{ meV}$. The next generation of cavity experiments in the frequency range (1–15) GHz address challenges related to sensitivity loss due to the reduction in volume¹⁶. In addition, the minimum detectable electromagnetic field in a cavity with a low noise linear amplifier is set by the standard quantum limit (SQL), an irreducible noise source that increases with frequency^{17,18}. Quantum-limited Josephson parametric amplifiers have been employed as preamplifiers in Dicke heterodyne receivers to readout haloscope cavities, as in ADMX and HAYSTAC experiments^{13,19}. The latter has recently circumvented the SQL with an electronic readout involving two JPAs to squeeze the receiver noise and improve the search speed by a factor of 2²⁰. Both the QUAX- $a\gamma$ ⁹ and QUAX- ae ²¹ haloscopes operated a JPA at approximately the SQL (system noise temperature of 0.9 K at $\sim 10 \text{ GHz}$), while in their future scientific runs they will take advantage of the large bandwidth of traveling wave parametric amplifiers^{22,23}.

The optimized dielectric resonator presented in this work, designed to resonate around 10 GHz with $\sim 500 \text{ MHz}$ tuning range, addresses the high frequency challenge in haloscope detection by exploiting the dielectric resonator concept^{24–28}. In this approach, the field profiles of higher-order resonant modes are properly modified by introducing dielectric structures that shape the cavity fields profiles, reducing the magnetic field amplitude on copper walls. This in turn helps to reduce dissipation and obtain higher quality factors. In addition, the dielectric material can suppress the out-of-phase components that negatively impact C_{mnl} when higher order modes are considered. The present dielectric resonator, shown in Fig. 1, has been improved compared to our previous realization²⁷ and we report unprecedentedly high quality factor values ever measured for a microwave resonator in multi-tesla fields.

II. HALOSCOPE CAVITY DESIGN

To optimize cavity design, we consider the cavity-based figure of merit $F = C_{mnl}^2 V^2 Q$. In fact, the appropriate expression to consider is that of the scan rate¹⁶, i.e. the speed at which haloscopes probe their targeted portions of axion parameter space for a fixed $g_{a\gamma\gamma}$ sensitivity

$$\frac{df}{dt} \approx \frac{g_{a\gamma\gamma}^4 \rho_a^2 B_0^4 C_{mnl}^2 V^2}{\Sigma^2 m_a^2 k_B^2 T_s^2} \left(\frac{\beta}{1+\beta} \right)^2 \frac{Q_c Q_a^2}{Q_c + Q_a}, \quad (3)$$

where T_s is the system noise temperature¹⁷. Note that this expression gives the nominal scan rate, and an efficiency factor should be added to account for dead time during data runs. RF cavities host transverse electric (TE_{mnl}) modes, transverse magnetic (TM_{mnl}) modes, and transverse electromagnetic (TEM_{mnl}) modes. If $\mathbf{B}_0 = B_0 \mathbf{z}$, the TE and TEM modes both lack axial E field components, resulting in null form factors. The modes that have non-zero form factors, and so of interest for axion detection, are the TM modes. For the TM_{010} mode the form factor is $C_{010} = 0.69$, while it is significantly smaller in higher order modes ($C_{020} = 0.13$ and $C_{030} = 0.05$).

An haloscope cavity must be operated under multi-tesla magnetic fields, that prevent the utilization of bulk niobium to reach high quality factors as is typically done in particle acceleration²⁹ and the conventional approach has been to use copper right circular cylinder cavities. A most promising technology to reach high Q resonators for haloscope search is based on utilization of films of high-magnetic-field tolerant, type II superconductors such as Nb-Ti, NbSn3, or ReBCO. Copper cavities sputtered with type II superconductors as NbTi and NbTiN films have been investigated at liquid helium temperature under DC magnetic fields up to 6 T, and the results of this search show that almost a factor 6 is lost at 6 T starting from $Q_0 = 1.4 \times 10^6$ measured at null field with a cavity resonant at 14.4 GHz, and rapidly decreasing for greater fields³⁰. Nonetheless, the $Q_0 B^2$ factor for this cavity at $B = 2 \text{ T}$ is still 5 times higher than in Cu, and it has been used to probe axions³¹, setting the limit $g_{a\gamma\gamma} < 1.03 \times 10^{-12} \text{ GeV}^{-1}$ for about $37 \mu\text{eV}$ mass (9.08 GHz). Low loss in 8 T magnetic field has also been accomplished with a biaxially textured YBCO cavity resonant at 6.9 GHz, with $Q_0 \sim 3 \times 10^5$ [32].

A strategy to probe for heavier axions is to increase V by employing multiple cell cavities to efficiently exploit the volume available in the bore of a superconducting solenoid³³, or to use long filter-based rectangular cavities when big volume dipolar magnets are available, as those that have been dismissed by LHC³⁴.

The highest frequency pathfinder run has been conducted using a copper cavity resonant at 26 GHz³⁵, even though in the 10-100 GHz frequency range the main experimental thrust is driven by the MADMAX project³⁶, in which a system of layers with alternating dielectric constants is used to obtain an enhanced axion-photon conversion. In this range a new class of thin-shell cavities has been recently proposed³⁷.

The dielectric cavity idea is based on dielectric materials properly placed inside traditional cylindrical resonant cavities,

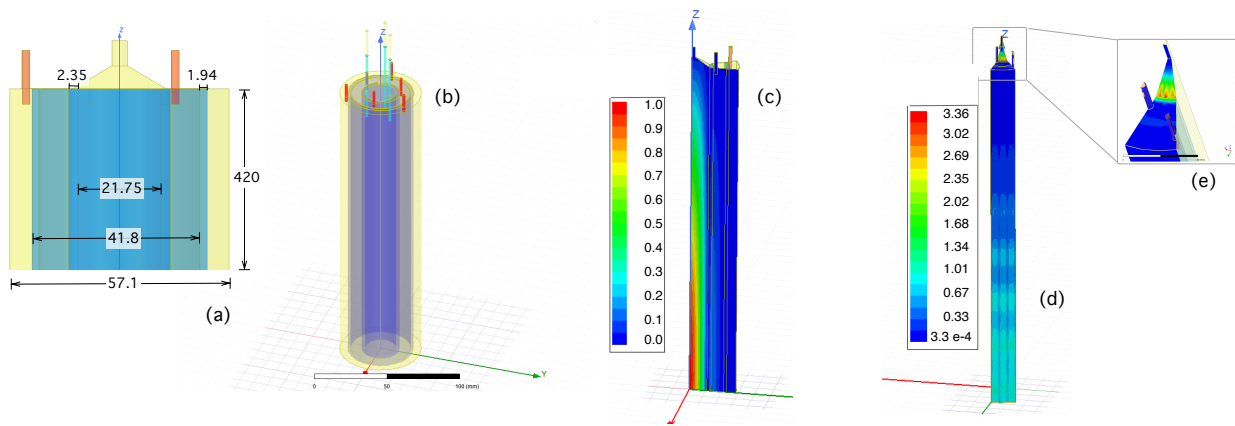


FIG. 1. (a) Cavity model (not in scale) and relevant dimensions. Tefflon pins are shown in orange. (b) 3D cavity model (in scale, half length) with copper external cylinder (light yellow), sapphire tubes (violet), tefflon pins (red) and sapphire rods employed for tuning the cavity frequency (light blue). (c) Electric field profile of the TM_{030} mode and (d) surface loss calculated via FEM simulations for a fraction of the overall cavity volume. (e) Surface loss in the cone structure.

operated in TM modes of higher order²⁴. Proof-of-concept experiments, including demonstration of a tuning mechanism, have been reported for a prototype Tefflon-wedge dielectric cavity at room temperature, following accurate design based on results from finite-element modeling²⁸. Highly pure aluminum oxide (99.7% Al_2O_3) was chosen as the dielectric material in a work by a different group, that investigated at 4 K the mode mixing of the TM_{030} mode with other cavity modes while tuning its frequency in the range (7.02-7.32) GHz²⁶.

Frequency tuning mechanisms in conventional experiments are based on transverse symmetry breaking by displacing one¹⁶ or multiple³⁸ metallic cylindrical rods. This can be performed by means of piezo-actuators or cryogenic motors. To demonstrate the resonator's effective tunability, careful comparison of simulation results and in situ tests are needed to address the problem of mode mixing, that takes place when other cavity modes cross with the axion-sensitive mode.

The proposed tuning methods for dielectric resonators encompass splitting the dielectric hollow cylinder vertically into two halves and moving them apart along the radial direction^{26,27}, or breaking the symmetry along the cavity axis by spitting the dielectric in the longitudinal direction³⁹.

In section VI we demonstrate a fine tuning mechanism based on dielectric rods, devised for preliminary tests of scientific runs. There is in fact much to be understood in the regime $Q_c > Q_a$, that this cavity allows to probe for the first time. The ~ 1 kHz-resolution, ~ 1 MHz-maximum range tuning we obtained reveal new challenges for data analysis that will be reported in a forthcoming work.

III. FEM ELECTROMAGNETIC SIMULATIONS AND METHODOLOGY

In our previous work²⁷ we reported tests conducted with a dielectric resonator composed of a couple of concentric hollow sapphire cylinders centered about the axis of a cylindrical copper cavity (see Fig. 1). In this type of cavity higher order

modes are exploited, and we focused on the TM_{030} mode for axion detection. The cavity was designed using finite element methods (FEM) so that the product CV was of the same order of magnitude for the TM_{030} mode in the dielectric cavity and the TM_{010} mode of a cylindrical copper cavity with same frequency and length. The realized cavity presented some practical issues, as the cylinders were not properly blocked, and therefore the cavity parameters were not reproducible at every cooling run. Most importantly, interference with a number of spurious modes was experimentally observed, that negatively impacted on the expected quality factor of the TM_{030} mode. Comparison of the results of an improved finite-element analysis of the adopted solution with in-situ tests, has shown that these additional modes arise from the circular grooves made in the copper endplates to hold the dielectric shells in place. Another important issue in the previous realization was related to the irregular shape of the sapphire tubes we used.

The cavity presented in this work has been re-designed to allow for centering of the dielectric shells axis to the copper cavity axis at cryogenic temperature. In the present configuration, the sapphire tubes are held in place by tefflon pins as shown in Fig. 1 (a) and (b), without the need of grooves in the copper endplates. The cylindrical body is 420 mm long and has 57 mm-diameter, while inner/external sapphire tubes have 21.75 mm/41.8 mm internal diameter and thickness 2.35 mm/1.94 mm. The cavity volume V is then 1.07761. In addition, the sapphire tubes we used for the present dielectric resonator have been shaped with mechanical tolerances chosen on the basis of FEM simulations. For instance, the required tolerances for the external/internal diameter of the internal cylinder were of ± 0.05 mm/ ± 0.15 mm. Tolerances on concentricity, linearity and coaxiality have also been given. Confinement of the electric field around the cavity axis, that is crucial to suppress dissipation in the cavity walls, is visualized in Fig. 1 (c).

The crystal tubes are grown using the Stephanov method by Rostox-N Ltd (Russia)⁴⁰. Pure sapphire raw materials (at least 99.99% Al_2O_3) are melted in a crucible made of Molybdenum

(2050 °C temperature) and the melt is then pulled through a ring-shaped Molybdenum component to obtain a sapphire tube.

To understand how each cavity component contribute to the overall losses and in turn influence the quality factor in this type of resonator, we have conducted FEM simulations taking into account the temperature dependence of losses in sapphire and copper. We have then measured Q_0 and f for several values of temperature in the range (4-293) K, and compared the experimental data with results from the simulations.

In table I we report the simulations parameters and the results, in which we decouple the losses in copper and sapphire, together with the corresponding experimental data.

The losses are additive, and their effect on the quality factor can be estimated through the simulations according to

$$\frac{1}{Q_0} = \frac{1}{Q_{Cu}} + \frac{1}{Q_{Sapph}} = R_s/G + \tan \delta, \quad (4)$$

where G is a constant that can be calculated depending on the cavity geometry and the chosen mode, R_s is the copper surface resistance and $\tan \delta$ is the tangent loss of sapphire. Here, Q_{Cu} and Q_{Sapph} are the quality factors of the modeled cavities with losses taking place entirely in copper or in sapphire, respectively.

The loss tangent at fixed temperature can significantly vary depending on crystallographic orientation, and the concentration of impurities or dislocations. A tangent loss value as low as 10^{-10} has been reported in sapphire at cryogenic temperatures⁴¹. In the simulations we used the values reported in table I, taken from Ref. 42, corrected to account for the different frequency⁴³.

In Fig. 1 (d) we show the surface loss profile of the TM_{030} in the copper walls, obtained running the FEM cavity model at 4.2 K. Cone structures at the end caps are used to reduce current dissipation, a critical aspect when highest quality factors are targeted. For instance, dissipation in the endcaps (shown in Fig. 1 (e)) for the mode of interest is limited to 16% of the overall dissipation due to copper finite conductivity.

As concerns the teflon screws, their contribution to overall cavity loss can be neglected. In fact, by attributing a value of $\tan \delta = 10^{-3}$ to teflon at room temperature the value $Q_0 = 5.8 \times 10^8$ is obtained in the modeled cavity with lossless metallic walls and dielectric cylinders.

The values of surface effective conductivity of copper σ_{Cu}^{eff} used in the model have been computed at 10.4 GHz in the anomalous skin depth regime⁴⁴.

For each temperature we account for the changes of cavity geometry using the thermal expansion coefficients of copper⁴⁵ and sapphire⁴⁶. We also considered the temperature dependence of the dielectric permittivity ϵ_r as reported in Table I, where the values in the third column have been calculated starting from the value we measured at 4.2 K ($\epsilon_r = 11.2$) and rescaling the data reported in Ref. 47. In table II we report simulation results for the modeled cavity at 4 K with a few values of ϵ_r including the values expected at 293 K and 4.2 K. As can be inferred by comparison of the simulation results in Table I and in Table II, we can ascribe a 19 MHz frequency shift when the cavity is cooled from room temperature to 4.2 K. An

additional shift of 9 MHz is instead determined by variation of the cavity geometry due to thermal expansion of the copper and sapphire cylinders. Moreover, simulation results in Table II show that the quality factor is not significantly affected by ϵ_r temperature changes.

This resonator is expected to have a quality factor as high as 11×10^6 at liquid helium temperature, calculated with parameters $\tan \delta = 1.19 \times 10^{-9}$, $\epsilon_r = 11.2$ and surface copper resistivity 0.00483Ω . The form factor, which can only be calculated through FEM modeling, is fixed to $C_{030} = 0.032$ as shown in table I and II. The $C_{030}V$ product is thus 34.6 cm^3 , compared to 164.6 cm^3 for the TM_{010} mode of an empty cylindrical cavity with the same length resonating at 10.4 GHz.

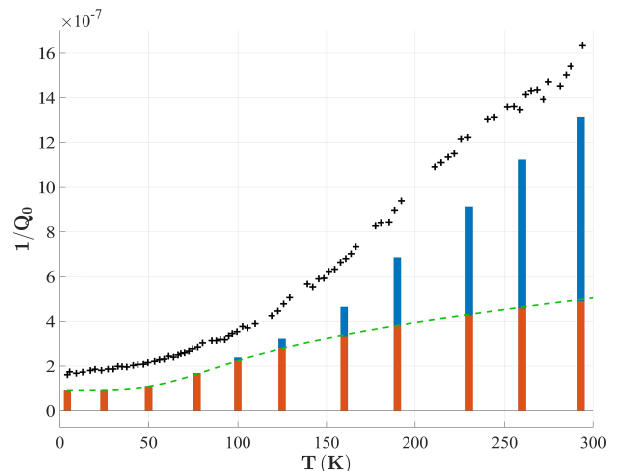


FIG. 2. Cavity dissipation according to the simulation. Loss in copper is represented in orange, that for sapphire is in blue. For comparison, we report in the same plot experimental data.

In Fig. 2 the simulation results show that at room temperature the losses in sapphire exceed those in copper, but they quickly decrease due to the loss tangent rapid drop. The quality factor of the modeled cavity saturates to 11 millions once below about 25 K, as at this point the copper conductivity does not vary significantly.

IV. MEASUREMENTS IN THE (4.2-293) K TEMPERATURE RANGE

We measure the resonant frequency of the TM_{030} mode and its quality factor by recording at a Vector Network Analyzer (VNA) the S_{21} scattering parameter of two coaxial antennas weakly coupled to the cavity mode.

The cavity can be cooled to ~ 4 K within a vacuum chamber designed to allow operation inside cryogenic dewars, and a thermometer calibrated to within 1 K has been used to measure the temperature. To ensure proper cooling of the cavity components, the chamber is loaded with He exchange gas. Fig. 3 shows the unloaded quality factor measured at different temperatures in the range (4.2-293) K, together with simulations results. Deviations from the data trend are observed

T (K)	σ_{Cu}^{eff}	ϵ_r	$\tan \delta$	f (GHz)	C_{030}	Q_{Sapph}	Q_{Cu}	Q_0	$Q_{0,exp}$
293	37.75	11.44	$4.87 \cdot 10^{-6}$	10.4468	0.0321	$1.21 \cdot 10^6$	$2.04 \cdot 10^6$	$7.60 \cdot 10^5$	$6.17 \cdot 10^5$
260	40.55	11.3889	$3.93 \cdot 10^{-6}$	10.4529	0.0321	$1.50 \cdot 10^6$	$2.18 \cdot 10^6$	$8.88 \cdot 10^5$	$7.28 \cdot 10^5$
230	43.70	11.3486	$2.88 \cdot 10^{-6}$	10.4576	0.0321	$2.05 \cdot 10^6$	$2.35 \cdot 10^6$	$1.09 \cdot 10^6$	$8.08 \cdot 10^5$
190	49.34	11.2953	$1.81 \cdot 10^{-6}$	10.4638	0.0320	$3.26 \cdot 10^6$	$2.64 \cdot 10^6$	$1.46 \cdot 10^6$	$1.09 \cdot 10^6$
160	55.53	11.2602	$7.84 \cdot 10^{-7}$	10.4676	0.0320	$7.51 \cdot 10^6$	$3.02 \cdot 10^6$	$2.12 \cdot 10^6$	$1.48 \cdot 10^6$
125	67.57	11.2288	$2.56 \cdot 10^{-7}$	10.4711	0.0320	$2.29 \cdot 10^7$	$3.58 \cdot 10^6$	$3.09 \cdot 10^6$	$2.19 \cdot 10^6$
100	84.17	11.2125	$8.87 \cdot 10^{-8}$	10.4729	0.0320	$6.62 \cdot 10^7$	$4.47 \cdot 10^6$	$4.15 \cdot 10^6$	$2.80 \cdot 10^6$
77	114.73	11.2044	$2.48 \cdot 10^{-8}$	10.4738	0.0320	$2.37 \cdot 10^8$	$6.10 \cdot 10^6$	$5.88 \cdot 10^6$	$3.33 \cdot 10^6$
50	174.57	11.2006	$7.52 \cdot 10^{-9}$	10.4744	0.0320	$7.82 \cdot 10^8$	$9.29 \cdot 10^6$	$9.02 \cdot 10^6$	$4.63 \cdot 10^6$
25	205.16	11.2002	$5.34 \cdot 10^{-9}$	10.4745	0.0321	$1.10 \cdot 10^9$	$1.09 \cdot 10^7$	$1.06 \cdot 10^7$	$5.48 \cdot 10^6$
4.2	207.04	11.2	$1.19 \cdot 10^{-9}$	10.4745	0.0321	$1.03 \cdot 10^9$	$1.10 \cdot 10^7$	$1.08 \cdot 10^7$	$6.23 \cdot 10^6$

TABLE I. Parameters set in the FEM simulations and results. Values are given at different temperatures for the effective copper conductivity σ_{Cu} , the relative dielectric constant ϵ_r and the tangent loss of sapphire. Losses in copper (Q_{Cu}) and in sapphire (Q_{Sapph}) are quantified by respectively running a model in which $\tan \delta$ is set to null value, and another model with lossless cavity walls. In the last column we report for comparison the measured unloaded quality factor $Q_{0,exp}$.

ϵ_r	f (GHz)	C_{030}	Q_{Cu}	Q_{Sapph}	Q_0
11.44	10.4556	0.0324	$1.14 \cdot 10^7$	$4.98 \cdot 10^9$	$1.11 \cdot 10^7$
11.3	10.4666	0.0322	$1.10 \cdot 10^7$	$4.96 \cdot 10^9$	$1.10 \cdot 10^7$
11.2	10.4745	0.0321	$1.09 \cdot 10^7$	$1.03 \cdot 10^9$	$1.08 \cdot 10^7$
11.1	10.4825	0.0319	$1.07 \cdot 10^7$	$1.21 \cdot 10^9$	$1.06 \cdot 10^7$

TABLE II. Simulation of the cavity at 4.2 K for different values of ϵ_r . The cavity frequency changes by 19 MHz (relative change) using the values reported in table I at room temperature and 4.2 K, while Q_0 and f are not significantly influenced.

when the TM_{030} mode mixes with other TE or TEM cavity modes.

While discrepancy between experimental data and the simulation results is within 20 % at room temperature, we observe in Fig. 3 an increasing deviation when the cavity is cooled down, suggesting that the dielectric losses in sapphire play a role also at low temperature. As shown in the inset of Fig. 3 the overall frequency shift of about 28 MHz predicted by the simulations for the temperature change from 293 K to 4.2 K, is in good agreement with the experimental results.

V. MEASUREMENTS AT LIQUID HELIUM TEMPERATURE IN HIGH MAGNETIC FIELDS

As this resonator is targeted to axion detection, it is important to measure its parameters under intense magnetic fields. Bulk copper undergoes magnetoresistance⁴⁸, but as shown in recent studies⁴⁹, the fractional change in quality factor of copper resonators immersed in B fields between 6 to 10 T is always positive, of about 2 % percent. Fig. 4 shows the results

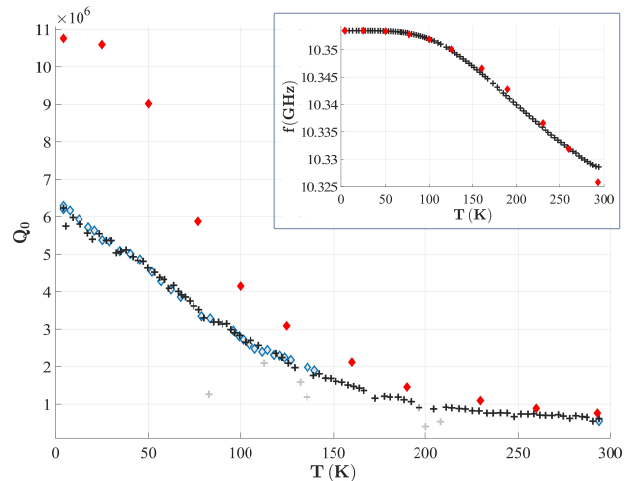


FIG. 3. Temperature dependence of the unloaded cavity quality factor. The light blue circles and black cross correspond to data series acquired during two different cooling runs. Red diamond data points are the values of quality factors obtained from the simulation. Inset: Temperature dependence of the TM_{030} mode frequency. Note that the resonant frequency values obtained from the simulations have been shifted by 121 MHz to match the value measured at 4 K.

of measurements with the present dielectric cavity inserted inside the 150 mm-diameter bore of an 8.1 T superconducting (SC) magnet of length 500 mm.

The magnet, manufactured by Cryogenic Ltd⁵⁰, was cooled in LHe and can be ramped up to a maximum of 8 T over less than 1 hour. To record the quality factor vs B-field dependence, we manually set current values using smaller steps between 0 and 2 T, and larger ones afterwards. Different regimes can be identified in the measured quality factor vs B-field plot: a rapid increase up to 0.8 T, a knee between 0.8 and 2 T fol-

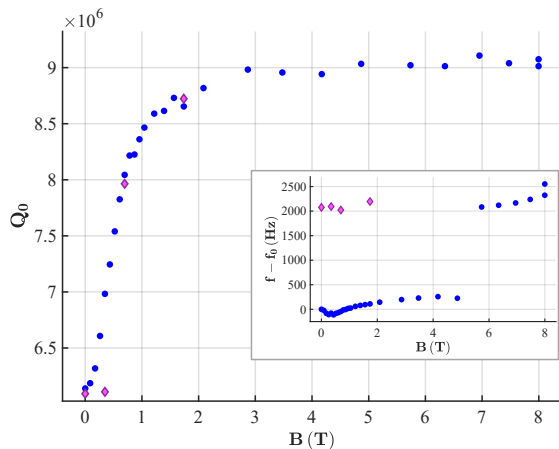


FIG. 4. Unloaded quality factor vs applied magnetic field. Diamond data points were acquired during magnetic field ramp down. Inset: Corresponding frequency shift. Note that at about 5 T the cavity frequency shifts by about 1.7 kHz, value comparable to the unloaded cavity linewidth $\Delta f = f/Q_0 \simeq 1.1$ kHz.

lowed by a small growth up to approximately 5 T. For values exceeding 5 T, a quality factor of about 9 millions is recorded, 50% greater than the value at null field. Remarkably, the measured maximum value at 8 T is only about 20% lower than the value expected from the simulations. The unexpected increase of quality factor in the presence of intense magnetic fields can be attributed to the magnetic properties of the dielectric cylinders. In fact, high purity sapphire crystals are known to host a number of paramagnetic impurities^{51,52} as Cr^{3+} , Fe^{3+} and Mo^{3+} . V^{2+} spin ensembles are also expected if the crystal has been subjected to radiation treatment or annealing procedures⁵¹ to relax internal stresses. Note that the present sapphire tubes have been annealed in vacuum at a temperature of 1900 C degrees.

Even when their concentration is as low as parts per billion, impurities can be detected and investigated by identifying avoided crossings between their electron spin resonances (ESR) and photonic whispering gallery modes^{51,52}. Measured properties are the zero field splitting, and the Landé g factors are obtained through the gradient of the frequency dependence on magnetic field of detected ESR transitions $df/dB = g\mu_B/\hbar$, with μ_B being the Bohr magneton. For instance, reported spin transition lines at zero field for Fe^{3+} , Cr^{3+} are at 12.03 GHz and 11.45 GHz respectively, while 8 lines in the range 8.7 – 10.4 GHz are attributed to the ion V^{2+} . With the application of a magnetic field, these absorption lines are tuned away from the TM_{030} cavity mode at the rate ~ 28 GHz/T for Cr and V impurities, and even faster for some transitions in Fe^{51} . Therefore we might expect that already at a few Tesla magnetic field amplitude, the related dissipative channel is significantly suppressed, giving a plausible explanation of the observed Q_0 vs B data trend.

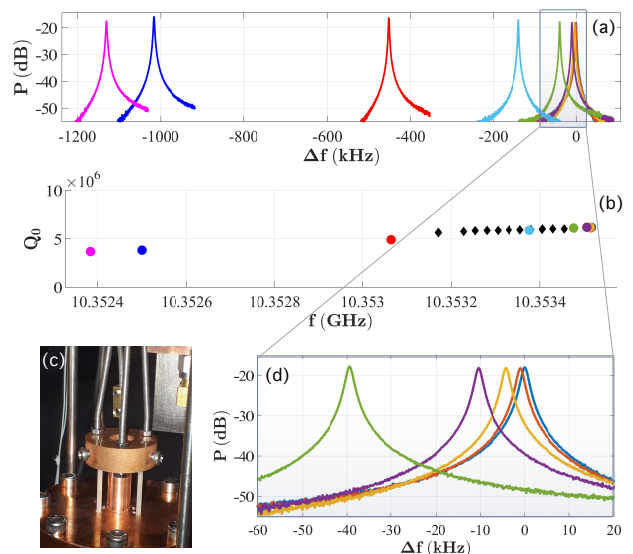


FIG. 5. Cryogenic frequency tuning. (a) Maximum 1.2 MHz frequency shift obtained and (b) measured Q_0 values exceeding Q_a . (c) Picture of the rods positioning system (top). Tuning is accomplished by displacing triplets of 2 mm-diameter sapphire rods relative to the top and bottom cavity endcaps. The triplets are anchored to copper blocks hanging on steel wires, that allow for controlling the rods position inside the cavity with an external micrometer. (d) Fine tuning. 1 kHz is the shown minimum frequency shift.

VI. TUNING

To make this resonator suitable for axion searches, the internal sapphire tube is substituted by two hollow half cylinders that allow for tuning the TM_{030} frequency up to about 500 MHz when moved apart along the radial direction^{26,27}. To demonstrate such a wide tuning range, a thorough study is under way involving simulations and in situ tests, here we report a tuning method that was implemented to tune over a narrow range the same cavity used in the measurements presented in previous sections. The objective in this case is to investigate the experimental requirements that are needed for running the experiment with a cavity exceeding the axion quality factor Q_a . In this study, that is currently ongoing in our laboratory, capability to shift the cavity central frequency within a fraction of its bandwidth is required. By displacing triplets of 2 mm-diameter sapphire rods relative to the top and bottom cavity endcaps, we shift the TM_{030} mode frequency without impacting the quality factor as shown in Fig. 5. The maximum shift we obtain is of about 1.2 MHz, with 1 kHz minimum frequency change, comparable with the cavity bandwidth.

VII. CONCLUSIONS

Microwave cavity resonators employed in axion dark matter search shall exhibit high quality factors in the presence of multi-tesla fields. In this work we have demonstrated the highest quality factor ever reported in the literature for a mi-

crowave cavity under intense magnetic fields, reaching more than 9 millions at 8 T. This result has been obtained with a copper cavity hosting two dielectric hollow cylinders, a design inspired by the idea that higher order modes in cavity resonators can effectively be used to address the high-frequency challenge in haloscope search when a low loss dielectric material is used to properly shape the electric field of the chosen mode. Practical limitations related to operating such a complex geometry at cryogenic temperatures²⁷ have been overcome in the present design, which has been perfected via comparison of the results of finite element modeling and tests in the laboratory. Measurements of the unloaded quality factor and resonance frequency of the TM₀₃₀ axion-sensitive mode have been conducted at various temperatures in the range (4.2 – 293) K, to determine the interplay between dissipation mechanisms in the copper walls and in the sapphire tubes. At 4.2 K we measured $Q_0 = 6.23 \times 10^6$, apparently limited by the copper finite conductivity as calculated taking into account the anomalous skin effect. Unexpectedly, the Q_0 measured under an 8 T-field exceeds the value measured at null field by about 50%. This experimental result is close to the values obtained from the simulations, in which we assumed that the tangent loss in sapphire decreases down to about 10^{-9} at 4.2 K, in agreement with several works reported in the literature.

As the losses in copper via magnetoresistance are expected to be within a few percent⁴⁹, in the measurements at low B-fields we identified a posteriori an additional loss channel due the magnetic susceptibility of sapphire. Paramagnetic impurities, previously investigated in high purity sapphire crystals, have indeed several zero point field transitions around the TM₀₃₀ frequency, that change at the rate $df/dB \sim 28$ GHz/T. Therefore at the high fields required for axion detection they are completely swept away from the TM₀₃₀ frequency.

As losses in copper can be accurately predicted, and having suppressed paramagnetic losses in sapphire at high magnetic field, we can obtain an estimate of the $\tan \delta$ of the sapphire tubes from the high field experimental results. The measured quality factor at 4.2 K is obtained in the FEM simulation if a tangent loss of 2.0×10^{-8} is assumed in place of 1.2×10^{-9} .

Our cavity development has involved precision metrology and modeling in simulation environments such as HFSS of ANSYS Electronics⁵³, demonstrating the capability to improve our understanding of this type of dielectric resonator and to optimize its future designs and tests that will include the wider tunability characterization. Across this study, measured and simulated results are in rather good agreement. Small quantitative differences are found within $\sim 1\%$, which is expected, given the complexity of the simulated model and taking into account mechanical tolerances.

We can use the scan rate expression given by eq. 3 to ponder the potential of this resonator for axion DM search with haloscopes. If β_{opt} is the value of coupling coefficient that maximizes the scan rate for a given Q_c/Q_a ratio⁸, we estimate about ~ 15 MHz/day for $\beta_{opt} \sim 7$ when a goal sensitivity to KSVZ family of models ($g_\gamma = -0.97$) at 95% confidence level ($\Sigma = 2$) is addressed, assuming quantum-limited readout and $B_0 = 14$ T. A reasonable signal integration time of about 50 s would then be needed for a single measurement. For compar-

ison, a conventional copper cavity resonating in its TM₀₁₀ at the same frequency ($r = 0.011$ m, $C_{010} = 0.5$, $Q_0 \simeq 20000$) would reach the same sensitivity at ~ 0.35 MHz/day at $\beta_{opt} \sim 4$, for which a feasible search is questionable.

ACKNOWLEDGMENTS

This material is based upon work supported by INFN (QUAX experiment) and the U.S. Department of Energy, Office of Science, National Quantum Information Science Research Centers, Superconducting Quantum Materials and Systems Center (SQMS) under the contract No. DE-AC02-07CH11359.

We are grateful to E. Berto (University of Padova and INFN) who substantially contributed to the mechanical realization of this cavity and of its tuning system. A. Benato (INFN) did part of the mechanical work and M. Zago (INFN) the mechanical drawings. The contribution of F. Calaon and M. Tessaro (INFN) to the experiment cryogenics and electronics is gratefully acknowledged. The cryogenic service of the Laboratori Nazionali di Legnaro provided the liquid helium to run the described experiments.

- ¹G. Bertone and D. Hooper, “History of dark matter,” *Rev. Mod. Phys.* **90**, 045002 (2018).
- ²R. D. Peccei and H. R. Quinn, “CP conservation in the presence of pseudoparticles,” *Phys. Rev. Lett.* **38**, 1440–1443 (1977).
- ³R. D. Peccei and H. R. Quinn, “Constraints Imposed by CP Conservation in the Presence of Instantons,” *Phys. Rev. D* **16**, 1791–1797 (1977).
- ⁴J. Ipser and P. Sikivie, “Are Galactic Halos Made of Axions?” *Phys. Rev. Lett.* **50**, 925 (1983).
- ⁵P. Sikivie, “Experimental Tests of the Invisible Axion,” *Phys. Rev. Lett.* **51**, 1415–1417 (1983), [Erratum: *Phys. Rev. Lett.* **52**, 695 (1984)].
- ⁶P. Sikivie, “Invisible axion search methods,” *Rev. Mod. Phys.* **93**, 015004 (2021).
- ⁷I. G. Irastorza and J. Redondo, “New experimental approaches in the search for axion-like particles,” *Progress in Particle and Nuclear Physics* **102**, 89–159 (2018).
- ⁸D. Kim, J. Jeong, S. Youn, Y. Kim, and Y. K. Semertzidis, “Revisiting the detection rate for axion haloscopes,” *JCAP* **03**, 066 (2020), arXiv:2001.05605 [hep-ex].
- ⁹D. Alesini, C. Braggio, G. Carugno, N. Crescini, D. D’Agostino, D. Di Gioacchino, R. Di Vora, P. Falferi, U. Gambardella, C. Gatti, G. Iannone, C. Ligi, A. Lombardi, G. Maccarrone, A. Ortolan, R. Pengo, A. Rettaroli, G. Ruoso, L. Taffarello, and S. Tocci, “Search for invisible axion dark matter of mass $m_a = 43 \mu\text{eV}$ with the quax- $a\gamma$ experiment,” *Phys. Rev. D* **103**, 102004 (2021).
- ¹⁰J. E. Kim, “Weak-interaction singlet and strong CP invariance,” *Phys. Rev. Lett.* **43**, 103–107 (1979).
- ¹¹M. Shifman, A. Vainshtein, and V. Zakharov, “Can confinement ensure natural cp invariance of strong interactions?” *Nuclear Physics B* **166**, 493–506 (1980).
- ¹²R. H. Dicke, “The measurement of thermal radiation at microwave frequencies,” *Review of Scientific Instruments* **17**, 268–275 (1946), <https://doi.org/10.1063/1.1770483>.
- ¹³T. e. a. Braine (ADMX Collaboration), “Extended search for the invisible axion with the axion dark matter experiment,” *Phys. Rev. Lett.* **124**, 101303 (2020).
- ¹⁴T. e. a. Braine (ADMX Collaboration), “Search for “invisible” axion dark matter in the 3.3–4.2 μeV mass range,” (2021), arXiv:2110.06096 [hep-ex].
- ¹⁵M. Dine, W. Fischler, and M. Srednicki, “A simple solution to the strong cp problem with a harmless axion,” *Physics Letters B* **104**, 199–202 (1981).
- ¹⁶I. Stern, A. A. Chisholm, J. Hoskins, P. Sikivie, N. S. Sullivan, D. B. Tanner, G. Carosi, and K. van Bibber, “Cavity design for high-frequency axion

- dark matter detectors,” *Review of Scientific Instruments* **86**, 123305 (2015), <https://doi.org/10.1063/1.4938164>.
- ¹⁷S. K. Lamoreaux, K. A. van Bibber, K. W. Lehnert, and G. Carosi, “Analysis of single-photon and linear amplifier detectors for microwave cavity dark matter axion searches,” *Phys. Rev. D* **88**, 035020 (2013).
- ¹⁸C. M. Caves, “Quantum limits on noise in linear amplifiers,” *Phys. Rev. D* **26**, 1817–1839 (1982).
- ¹⁹L. Zhong, S. Al Kenany, K. M. Backes, B. M. Brubaker, S. B. Cahn, G. Carosi, Y. V. Gurevich, W. F. Kindel, S. K. Lamoreaux, K. W. Lehnert, S. M. Lewis, M. Malnou, R. H. Maruyama, D. A. Palken, N. M. Rapidis, J. R. Root, M. Simanovskaia, T. M. Shokair, D. H. Speller, I. Urdinaran, and K. A. van Bibber, “Results from phase I of the haystac microwave cavity axion experiment,” *Phys. Rev. D* **97**, 092001 (2018).
- ²⁰K. M. Backes, D. A. Palken, S. A. Kenany, B. M. Brubaker, S. B. Cahn, A. Droster, G. C. Hilton, S. Ghosh, H. Jackson, S. K. Lamoreaux, A. F. Leder, K. W. Lehnert, S. M. Lewis, M. Malnou, R. H. Maruyama, N. M. Rapidis, M. Simanovskaia, S. Singh, D. H. Speller, I. Urdinaran, L. R. Vale, E. C. van Assendelft, K. van Bibber, and H. Wang, “A quantum enhanced search for dark matter axions,” *Nature* **590**, 238–242 (2021).
- ²¹N. Crescini, D. Alesini, C. Braggio, G. Carugno, D. D’Agostino, D. Di Gioacchino, P. Falferi, U. Gambardella, C. Gatti, G. Iannone, C. Ligi, A. Lombardi, A. Ortolan, R. Pengo, G. Ruoso, and L. Taffarello (QUAX Collaboration), “Axion search with a quantum-limited ferromagnetic haloscope,” *Phys. Rev. Lett.* **124**, 171801 (2020).
- ²²L. Planat, A. Ranadive, R. Dassonville, J. Puertas Martínez, S. Léger, C. Naud, O. Buisson, W. Hasch-Guichard, D. M. Basko, and N. Roch, “Photonic-crystal josephson traveling-wave parametric amplifier,” *Phys. Rev. X* **10**, 021021 (2020).
- ²³A. Ranadive, M. Esposito, L. Planat, E. Bonet, C. Naud, O. Buisson, W. Guichard, and N. Roch, “A reversed kerr traveling wave parametric amplifier,” (2021), [arXiv:2101.05815](https://arxiv.org/abs/2101.05815) [quant-ph].
- ²⁴D. E. Morris, “An electromagnetic detector for relic axions,” (1984).
- ²⁵B. T. McAllister, G. Flower, L. E. Tobar, and M. E. Tobar, “Tunable supermode dielectric resonators for axion dark-matter haloscopes,” *Phys. Rev. Applied* **9**, 014028 (2018).
- ²⁶J. Kim, S. Youn, J. Jeong, W. Chung, O. Kwon, and Y. K. Semertzidis, “Exploiting higher-order resonant modes for axion haloscopes,” *J. Phys. G* **47**, 035203 (2020), [arXiv:1910.00793](https://arxiv.org/abs/1910.00793) [physics.ins-det].
- ²⁷D. e. a. Alesini (QUAX), “Realization of a high quality factor resonator with hollow dielectric cylinders for axion searches,” *Nucl. Instrum. Meth. A* **985**, 164641 (2021), [arXiv:2004.02754](https://arxiv.org/abs/2004.02754) [physics.ins-det].
- ²⁸A. P. Quiskamp, B. T. McAllister, G. Rybka, and M. E. Tobar, “Dielectric-booster sensitivity to cylindrical azimuthally varying transverse-magnetic resonant modes in an axion haloscope,” *Phys. Rev. Applied* **14**, 044051 (2020).
- ²⁹H. S. Padamsee, “Superconducting radio-frequency cavities,” *Annual Review of Nuclear and Particle Science* **64**, 175–196 (2014), <https://doi.org/10.1146/annurev-nucl-102313-025612>.
- ³⁰D. Di Gioacchino, C. Gatti, D. Alesini, C. Ligi, S. Tocci, A. Rettaroli, G. Carugno, N. Crescini, G. Ruoso, C. Braggio, P. Falferi, C. S. Gallo, U. Gambardella, G. Iannone, G. Lamanna, A. Lombardi, R. Mezzena, A. Ortolan, R. Pengo, E. Silva, and N. Pompeo, “Microwave losses in a dc magnetic field in superconducting cavities for axion studies,” *IEEE Transactions on Applied Superconductivity* **29**, 1–5 (2019).
- ³¹D. Alesini *et al.*, “Galactic axions search with a superconducting resonant cavity,” *Phys. Rev. D* **99**, 101101 (2019), [arXiv:1903.06547](https://arxiv.org/abs/1903.06547) [physics.ins-det].
- ³²D. Ahn, O. Kwon, W. Chung, W. Jang, D. Lee, J. Lee, S. Woo Youn, H. Byun, D. Youm, and Y. K. Semertzidis, “First prototype of a biaxially textured $\text{YBa}_2\text{Cu}_3\text{O}_{7-x}$ microwave cavity in a high magnetic field for dark matter axion search,” (2021), [arXiv:2103.14515](https://arxiv.org/abs/2103.14515) [physics.ins-det].
- ³³J. Jeong, S. Youn, S. Bae, J. Kim, T. Seong, J. E. Kim, and Y. K. Semertzidis, “Search for invisible axion dark matter with a multiple-cell haloscope,” *Phys. Rev. Lett.* **125**, 221302 (2020).
- ³⁴A. A. Melcón *et al.* (CAST), “First results of the CAST-RADES haloscope search for axions at $34.67 \mu\text{eV}$,” (2021), [arXiv:2104.13798](https://arxiv.org/abs/2104.13798) [hep-ex].
- ³⁵B. T. McAllister, G. Flower, E. N. Ivanov, M. Goryachev, J. Bourhill, and M. E. Tobar, “The organ experiment: An axion haloscope above 15 ghz,” *Physics of the Dark Universe* **18**, 67–72 (2017).
- ³⁶B. Majorovits, “MADMAX: A new road to axion dark matter detection,” **1342**, 012098 (2020).
- ³⁷C.-L. Kuo, “Symmetrically Tuned Large-Volume Conic Shell-Cavities for Axion Searches,” *JCAP* **02**, 018 (2021), [arXiv:2010.04337](https://arxiv.org/abs/2010.04337) [physics.ins-det].
- ³⁸M. Simanovskaia, A. Droster, H. Jackson, I. Urdinaran, and K. van Bibber, “A symmetric multi-rod tunable microwave cavity for a microwave cavity dark matter axion search,” *Review of Scientific Instruments* **92**, 033305 (2021), <https://doi.org/10.1063/5.0016125>.
- ³⁹B. T. McAllister, G. Flower, L. E. Tobar, and M. E. Tobar, “Tunable Supermode Dielectric Resonators for Axion Dark-Matter Haloscopes,” *Phys. Rev. Applied* **9**, 014028 (2018), [arXiv:1705.06028](https://arxiv.org/abs/1705.06028) [physics.ins-det].
- ⁴⁰“Rostox-n Ltd, <https://rostox-n.com/en>,” .
- ⁴¹S. N. Buckley, P. Agnew, and G. P. Pells, “Cryogenic dielectric properties of sapphire at 2.45 GHz,” **27**, 2203–2209 (1994).
- ⁴²V. Braginsky, V. Ilchenko, and K. S. Bagdassarov, “Experimental observation of fundamental microwave absorption in high-quality dielectric crystals,” *Physics Letters A* **120**, 300–305 (1987).
- ⁴³N. A. Shtin, J. M. L. Romero, and E. Prokhorov, “Theory of fundamental microwave absorption in sapphire ($\alpha\text{-al } 2 \text{ o } 3$),” *Journal of Applied Physics* **106**, 104115 (2009).
- ⁴⁴S. Calatroni, “A Mathematica Notebook for the calculation of the anomalous skin effect in copper,” (2020).
- ⁴⁵N. Simon, E. Drexler, and R. Reed, “Nist monograph 177 properties of copper and copper alloys at cryogenic temperature,” US Government printing Office , 2–27 (1992).
- ⁴⁶P. E. Bradley, R. Radebaugh, *et al.*, “Properties of selected materials at cryogenic temperatures,” *NIST Publ* **680**, 1–14 (2013).
- ⁴⁷J. Krupka, K. Derzakowski, M. Tobar, J. Hartnett, and R. G. Geyer, “Complex permittivity of some ultralow loss dielectric crystals at cryogenic temperatures,” *Measurement Science and Technology* **10**, 387 (1999).
- ⁴⁸J. de Launay, R. Dolecek, and R. Webber, “Magnetoresistance of copper,” *Journal of Physics and Chemistry of Solids* **11**, 37–42 (1959).
- ⁴⁹S. Ahn, S. Youn, J. Yoo, D. Kim, J. Jeong, M. Ahn, J. Kim, D. Lee, J. Lee, T. Seong, and Y. Semertzidis, “Magnetoresistance in copper at high frequency and high magnetic fields,” *Journal of Instrumentation* **12**, P10023–P10023 (2017).
- ⁵⁰<https://cryogenic.co.uk/>.
- ⁵¹W. G. Farr, D. L. Creedon, M. Goryachev, K. Benmessai, and M. E. Tobar, “Ultrasensitive microwave spectroscopy of paramagnetic impurities in sapphire crystals at millikelvin temperatures,” *Phys. Rev. B* **88**, 224426 (2013).
- ⁵²V. Giordano, C. Fluhr, and B. Dubois, “Magnetic sensitivity of the microwave cryogenic sapphire oscillator,” *Journal of Applied Physics* **127**, 184101 (2020), <https://doi.org/10.1063/5.0007131>.
- ⁵³<https://www.ansys.com/products/electronics/ansys-hfss>.
- ⁵⁴S. Al Kenany, M. Anil, K. Backes, B. Brubaker, S. Cahn, G. Carosi, Y. Gurevich, W. Kindel, S. Lamoreaux, K. Lehnert, S. Lewis, M. Malnou, D. Palken, N. Rapidis, J. Root, M. Simanovskaia, T. Shokair, I. Urdinaran, K. van Bibber, and L. Zhong, “Design and operational experience of a microwave cavity axion detector for the 20–100 μeV range,” *Nuclear Instruments and Methods in Physics Research Section A: Accelerators, Spectrometers, Detectors and Associated Equipment* **854**, 11–24 (2017).