Sub-GeV Atmospheric Neutrinos and CP-Violation in DUNE

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We propose to use the unique event topology and reconstruction capabilities of liquid argon time projection chambers to study sub-GeV atmospheric neutrinos. The detection of low energy recoiled protons in DUNE allows for a determination of the leptonic CP-violating phase independent from the accelerator neutrino measurement. Our findings indicate that this analysis can exclude several values of $\delta_{CP}$ beyond the 3$\sigma$ level. Moreover, the determination of the sub-GeV atmospheric neutrino flux will have important consequences in the detection of diffuse supernova neutrinos and in dark matter experiments.

I. INTRODUCTION

Atmospheric neutrinos, produced by cosmic-ray interactions in the Earth’s atmosphere, have played a crucial role in the discovery of neutrino oscillations [1], the only evidence of non-zero neutrino masses [2–3]. Even now, atmospheric neutrinos contribute significantly to our understanding of neutrino oscillations and mixing in the lepton sector [4–5]. In this Letter, we are particularly interested in such neutrinos with energies in the 100 MeV to 1 GeV region, i.e. sub-GeV atmospheric neutrinos. The oscillation phenomenology of this sample is exceptionally rich [6–15]. The physical reason behind this is twofold. First, for baselines comparable to the Earth’s radius, oscillation of sub-GeV neutrinos are strongly affected by both solar and atmospheric mass splittings. Second, the broad energy spectrum and large matter effects induced both solar and atmospheric mass splittings. Second, the oscillations in sub-GeV atmospheric neutrinos to probe CP-violating effects. All this results in very poor reconstruction of the neutrino energy, direction and flavor (between $\nu_e$ and $\bar{\nu}_e$), making the use of sub-GeV atmospheric neutrinos to probe CP-violation in Cherenkov detectors impractical, unless detectors are gigantic, at the multi-megaton scale [20].

In liquid argon time projection chambers (LArTPCs), the situation is completely different. The LArTPC technology allows for excellent reconstruction of neutrino event topologies by detecting the tracks of all charged particles and identifying them by topology and energy loss. Recently it was shown that protons with kinetic energy above 21 MeV can be efficiently identified in the ArgoNeut experiment [21]. Their three-momenta can still be reconstructed with good resolution, which will allow for a pioneering measurement of sub-GeV neutrino energies and angles. Besides, the capability of detecting these protons allows for statistical separation between sub-GeV neutrinos and antineutrinos, since the former is significantly more likely to kick out a single proton from Argon than the latter [21]. Together with the fact that the cross section for neutrinos is about a factor of 2 larger than the latter [21]. Together with the fact that the cross section for neutrinos is about a factor of 2 larger than the latter. However, protons with less than about 1.4 GeV of kinetic energy do not emit any Cherenkov light in water, and thus are as invisible as neutrons in these detectors.

In neutrino scattering events, the kinematics of the outgoing lepton bears correlation with the neutrino energy and direction. Nevertheless, at these low energies much of such correlation is lost, and there is a large spread in outgoing lepton momentum and angle. On top of that, the lack of charge identification impedes the separation of events originated from neutrinos or antineutrinos, concealing CP-violating effects. All this results in very poor reconstruction of the neutrino energy, direction and flavor (between $\nu_e$ and $\bar{\nu}_e$), making the use of sub-GeV atmospheric neutrinos to probe CP-violation in Cherenkov detectors impractical, unless detectors are gigantic, at the multi-megaton scale [20].

At present, only Cherenkov light detectors like Super-Kamiokande and IceCube are large enough to have significant sensitivity to the broad spectrum of atmospheric neutrinos. Nevertheless, a precise measurement of sub-GeV neutrinos is still lacking. The difficulty in studying these neutrinos is related to the event reconstruction which is very challenging at these low energies. When a sub-GeV neutrino scatters on a nucleon via a charged current interaction, it produces a charged lepton and recoils the nucleon isospin partner, for instance $\nu_e n \rightarrow e^- p^+$. However, protons with less than about 1.4 GeV of kinetic energy do not emit any Cherenkov light in water, and thus are as invisible as neutrons in these detectors.
(DUNE) \cite{22} will be able to measure sub-GeV atmospheric neutrinos and extract information on $\delta_{\text{CP}}$ complementary to the accelerator neutrino program. In addition to $\delta_{\text{CP}}$, the study of sub-GeV atmospheric neutrinos will have major impact in the determination of the diffuse supernova neutrino background and the neutrino background in dark matter direct detection experiments, as well as in searches for new physics in the neutrino sector.

II. PHYSICS WITH LOW-ENERGY ATMOSPHERIC NEUTRINOS

In general terms, neutrino oscillations are driven by a phase $\propto (\Delta m_{ij}^2/\text{eV}^2)(L/\text{km})(\text{GeV}/E)$, where $L$ is the distance traveled between neutrino production and detection, $E$ is the neutrino energy, and $\Delta m_{ij}^2 = m_i^2 - m_j^2$ is the squared mass splitting. When $E \gtrsim 1$ GeV, oscillations are induced largely by the aptly-named atmospheric mass splitting $|\Delta m_{31}^2| \approx 2.5 \times 10^{-3}$ eV$^2$ \cite{23,24}, and they develop on scales $L \sim \mathcal{O}(R_E)$, the radius of the Earth. Oscillations of atmospheric neutrinos with energies $100 \text{ MeV} < E < 1$ GeV, are governed by both the atmospheric mass splitting and the smaller solar mass splitting, $\Delta m_{21}^2 \approx 7.4 \times 10^{-5}$ eV$^2$ \cite{23,30}. In what follows, we will consider two aspects of major significance to our analysis, $CP$ violation and matter effects. We adopt the usual parametrization for neutrino mixing \cite{31}.

To set convention, we define the zenith angle such that $\cos \theta_z = -1$ corresponds to neutrinos coming from directly below the detector, while $\cos \theta_z = 0$ indicates the horizon direction.

First we discuss the effects of $\delta_{\text{CP}}$ in oscillations of sub-GeV neutrinos. In vacuum, for simplicity, the $CP$-violating term in neutrino oscillation probability is given by \cite{32}

$$P_{\text{CP}} = -8 J_s \sin \delta_{\text{CP}} \sin \Delta_{21} \sin \Delta_{31} \sin \Delta_{32},$$

which includes the Jarlskog invariant \cite{33,34} $J_s \sin \delta_{\text{CP}}$ (in our convention) and $\Delta_{ij} \equiv \Delta m_{ij}^2 L/4E$ are the oscillation phases. Oscillations of beam neutrinos probe the atmospheric splitting $\Delta_{31} \sim \mathcal{O}(1)$, while $\Delta_{21} \ll 1$. There, the $CP$ term is suppressed by $\Delta m_{21}^2/\Delta m_{31}^2 \times \pi/2 \sim 1/20$ due to the fact that oscillations driven by $\Delta m_{21}^2$ do not have time to develop. This yields $P_{\text{CP}} \approx -0.4 J_s \sin \delta_{\text{CP}} \sin \Delta_{31} \sin \Delta_{32}$. Sub-GeV atmospheric neutrino oscillations, on the other hand, probe the solar splitting. In this case, the oscillations driven by $\Delta m_{32,31}^2$ are fast and average out. The resulting factor is just $1/2$, leading to a much larger $CP$-violating term relative to beam neutrinos, namely $P_{\text{CP}} \approx -4 J_s \sin \delta_{\text{CP}} \sin \Delta_{21}$ with $\Delta_{21} \sim \mathcal{O}(1)$.

In Fig. \ref{fig:osc-prob} we present several oscillation probability curves\cite{4} as function of neutrino energy for $\nu_e \rightarrow \nu_e$ (black), $\nu_\mu \rightarrow \nu_e$ (red), $\nu_\mu \rightarrow \nu_\mu$ (blue), and $\nu_\tau \rightarrow \nu_e$ (dashed red) for various values of the zenith angle $\cos \theta_z$ as indicated, and $\delta_{\text{CP}} = 0$ (top panel) or $\delta_{\text{CP}} = 3\pi/2$ (all other panels). Earth’s matter profile was implemented using the PREM model \cite{35}. In the upper panel, the red and blue lines lie on top of each other.

$\nu_e \rightarrow \nu_\mu$ (blue) and $\nu_\mu \rightarrow \nu_e$ (dashed red), various zenith angles and $\delta_{\text{CP}} = 0$ (upper panel) or $\delta_{\text{CP}} = 3\pi/2$ (all other panels). Focusing on the blue and red curves in the first two panels, we observe a large effect of $\delta_{\text{CP}}$, as indicated, and $\delta_{\text{CP}} = 0$ (top panel) or $\delta_{\text{CP}} = 3\pi/2$ (all other panels). Earth’s matter profile was implemented using the PREM model \cite{35}. In the upper panel, the red and blue lines lie on top of each other.

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The second crucial feature of sub-GeV atmospheric

\begin{figure}[h]
\centering
\includegraphics[width=\textwidth]{osc-prob.png}
\caption{Oscillation probabilities for $\nu_e \rightarrow \nu_e$ (black), $\nu_\mu \rightarrow \nu_e$ (red), $\nu_\mu \rightarrow \nu_\mu$ (blue), and $\nu_\tau \rightarrow \nu_e$ (dashed red) for various values of the zenith angle $\cos \theta_z$ as indicated, and $\delta_{\text{CP}} = 0$ (top panel) or $\delta_{\text{CP}} = 3\pi/2$ (all other panels). Earth’s matter profile was implemented using the PREM model \cite{35}. In the upper panel, the red and blue lines lie on top of each other.}
\end{figure}

\begin{enumerate}
\item\footnote{See \url{https://imgur.com/HowUn1w} for an animation of how $\delta_{\text{CP}}$ changes oscillation probabilities as a function of zenith angle and neutrino energy.}
\end{enumerate}
neutrino oscillations are matter effects. Interactions with matter in the Earth, specifically in the dense mantle and core, may significantly modify neutrino oscillations. These effects are quite rich and have been studied in great depth \cite{10,11}. Here we restrict ourselves to review some oscillation aspects and provide a few examples.

Up-going atmospheric neutrinos that traverse the Earth may go through an MSW resonance \cite{16,17} in the solar sector, maximizing oscillations between $\nu_e$ and $\nu_{\mu, \tau}$, when

$$\Delta m^2_{21} \cos \theta_{12} = 2\sqrt{2}EG_F n_e,$$

(2)

where $G_F$ is the Fermi constant, and $n_e$ is the electron number density. In the solar sector, the MSW resonance happens only for neutrinos, not for antineutrinos, as observed in oscillation of neutrinos produced in the Sun. We will focus on the $\nu_e \rightarrow \nu_e$ oscillation dependence on the zenith angle, shown as black curves in the different panels of Fig. 1. In the crust (first or second panels, $-0.44 < \cos \theta_z$), mantle (third and fourth panels, $-0.84 < \cos \theta_z < -0.44$) and core (bottom panel, $\cos \theta_z < -0.84$), the MSW resonant energies are found to be around 180, 130, and 50 MeV, respectively. Although this energy in Earth’s core is below 100 MeV, another type of resonance occurs about $E \sim 170$ MeV, a parametric resonance \cite{18,19,24}. A parametric resonance happens when changes to the matter density profile occur on the same scale as the neutrino oscillation length. The phenomenon is analogous to a resonant spring oscillator. The fast oscillations (which begin to slow down as $E \rightarrow 1$ GeV) are induced by the atmospheric mass splitting. Note that, due to the near-maximal value of $\theta_{23}$, $\nu_e$ oscillates approximately equally into $\nu_\mu$ and $\nu_\tau$.

The $CP$-violating and matter effects displayed in Fig. 1 show that the $\delta_{CP}$ effect is broad in neutrino energy, but there are large variations of oscillation curves for different zenith angles. Therefore, the precise reconstruction of the neutrino energy will not be as important as the determination of the incoming neutrino direction for the measurement of $\delta_{CP}$. \LArTPCs have excellent energy resolution and tracking reconstruction, and hence the incoming neutrino direction may be determined by considering the full event topology in charged current quasi-elastic events, $\nu_n \nu_n \rightarrow \ell^- p^+$. In the next Section, we will discuss the details of our simulation of sub-GeV atmospheric neutrinos and how we take the nuclear physics effects into account.

III. SIMULATION DETAILS

To simulate the atmospheric neutrino flux at sub-GeV energies, we use Ref. \cite{39}. The atmospheric neutrinos flux for a given flavor is parametrized by

$$\Phi_\alpha(E) = \Phi_{\alpha,0} f_\alpha(E) \left( \frac{E}{E_0} \right)^\delta,$$

(3)

where $f_\alpha(E)$ gives the shape of the neutrino energy spectrum for each flavor (obtained from simulations, see Ref. \cite{39}), $\Phi_{\alpha,0}$ is the normalization of the flavor $\alpha$ ($\nu_e, \nu_\mu, \nu_\tau, \bar{\nu}_\mu$), $E_0$ is an arbitrary reference energy, and $\delta$ accounts for spectral distortions. To account for unknowns on the meson production in the atmosphere, we consider systematic uncertainties on the following quantities: overall normalization (40%); the ratio $r_e$ between $\nu_e$ and $\nu_\mu$ fluxes (5%); the ratio $r_p$ between neutrinos and antineutrinos fluxes (2%); and the spectral distortion parameter $\delta$ with 0.2 absolute uncertainty.

Neutrino events in the DUNE detector will be classified by topology. We consider events with a charged lepton (electrons or muons) and up to 2 outgoing protons and pions, namely $CC-Np0\pi\pi (N = 0,1,2)$. The interaction of neutrinos scattering on argon was modeled with the NuWro event generator \cite{37}. This is an important step as recoiled nucleons may re-interact still inside the nucleus, a process typically referred to as final state interactions or intra-nuclear cascades. A pictorial representation of intra-nuclear cascades is shown in Fig. 2. To account for detector response, a cut on the minimum proton kinetic energy of 30 MeV was implemented \cite{22}. Momentum resolutions of 5%, 5% and 10% for electrons, muons and protons were assumed \cite{38} as well as conservative angular resolutions of $5^\circ, 5^\circ$ and $10^\circ$, respectively \cite{39}. We have checked that these angular and energy resolutions are not limiting factors in our results.

Incoming neutrino energy and direction were estimated by reconstructing the energy and direction of all outgoing charged particles, summing up their four momenta, and subtracting the four momenta of initial nucleons assuming they were at rest. For example, in a $CC-2p0\pi$ event the reconstructed neutrino energy would be given by $E_{\nu} = E_{\ell} + K_P^{(1)} + K_P^{(2)}$, where $K_P$ indicates the proton kinetic energy. Besides the imperfect detector response, intra-nuclear cascades as well as outgoing neutrinos (which we consider to always go undetected) can affect the neutrino energy and direction reconstruction. We find that the largest contribution to energy and direction mis-construction arrives from intra-nuclear cascades \cite{37}. A similar technique was proposed in Ref. \cite{40,41} to improve the DUNE sensitivity for dark matter annihilation in the Sun using pointing.

To evaluate the experimental sensitivity to $\delta_{CP}$, we have calculated the oscillation probabilities for $-1 \leq \cos \theta_z \leq 1$ and 100 MeV $\leq E_\nu \leq 1$ GeV, assuming the
The sensitivity to $\delta_{CP}$ for an input value of $\delta_{CP} = 3\pi/2$ is shown in Fig. 3 assuming a 400 kton-year exposure. The individual $\Delta \chi^2$ contribution for each topology is shown, as well as the combined fit. A remarkable sensitivity to $\delta_{CP}$ may be achieved, allowing for excluding regions of the parameter space beyond the $3\sigma$ level.

Several factors contribute to this sensitivity. As already discussed, the $CP$ violation effect for sub-GeV atmospheric neutrinos is a sizable effect, an order of magnitude larger than the corresponding one for beam neutrinos. To observe $CP$ violation, one should be able to independently measure oscillations of neutrinos and antineutrinos and/or the time-conjugated channels $\nu_\mu \rightarrow \nu_e$ and $\bar{\nu}_e \rightarrow \bar{\nu}_\mu$. The event topology reconstruction in LArTPCs allows for counting the number of protons in the final state. At these low energies, a neutrino interaction is more likely to kick out a proton from a nucleus than an antineutrino interaction, and vice-versa for neutrons – therefore, the CC-1p0$\pi_0$ topology is rich while CC-0p0$\pi_0$ is antineutrino rich. Combining these two samples allows for measuring, statistically, the flux of $\nu$ and $\bar{\nu}$ from the atmosphere. Besides, the incoming neutrino zenith direction has a typical spread, mainly due to intranuclear cascades, between $\Delta \theta \sim 20^\circ - 30^\circ$ using our reconstruction technique, except for the CC-0p0$\pi_0$ topology which has $\Delta \theta \sim 50^\circ$. This allows to infer the neutrino direction fairly well, disentangling the rich oscillation effects, discussed in Sec. 11 for different baselines. These aspects indicate a synergy between each distinct topology, as it can be seen in Fig. 3 the sum of the individual $\Delta \chi^2$ contributions for each topology is significantly below the combined sensitivity.

With respect to the systematics, we have found that DUNE constrains the pull parameters beyond the uncertainties adopted here, evidencing that the experimental sensitivity is not induced by any prior uncertainty on the atmospheric fluxes, and therefore is quite robust. We have found that the overall normalization of sub-GeV atmospheric neutrinos can be constrained at the 2% level, while $r_\nu$ and $r_{\bar{\nu}}$ are determined at the 2% and 1% level, respectively, and the spectral index of the flux is constrained as $\delta = \pm 0.02$. This is an important finding for the detection of the neutrino flux from diffuse supernovae, as low energy atmospheric neutrinos are among the largest backgrounds to this measurement [13]. Besides, low energy atmospheric neutrinos are an important component of the neutrino floor, the neutrino-nucleus coherent scattering background in dark matter direct detection [14].

The sensitivity to $\delta_{CP}$ obtained here, though not as powerful as the one obtained with beam neutrinos, is quite competitive. This will provide a very important cross check for the determination of $\delta_{CP}$ involving energies and baselines very different from those in beam neutrinos. The experimental sensitivity to $\delta_{CP}$ is limited by statistics and the reconstruction of the original neutrino zenith angle. We expect that the addition of multi-GeV atmospheric neutrinos will further improve the $CP$ sensitivity, but we leave this study for future work. Moreover, this distinct configuration will also boost new physics searches such as effective non-standard neutrino interactions and sterile neutrino scenarios.

As a side remark, since the MSW resonance is related to $\Delta m_{23}^2$, and oscillations here are driven by $\theta_{12}$ and $\theta_{23}$, this sample may also be used to constrain other oscillation parameters. In Fig. 4 the experimental sensitivity

![Sub−GeV Atmospheric Neutrinos](image)

**FIG. 3.** DUNE sensitivity to the leptonic CP violating phase $\delta_{CP}$ using sub-GeV atmospheric neutrinos, for an input value $\delta_{CP} = 3\pi/2$. An exposure of 400 kton-year was assumed.
in the $\delta_{CP} \times \sin^2 \theta_{23}$ plane (again, after marginalizing over systematic uncertainties) is shown for DUNE sub-GeV atmospheric neutrino analysis, again for 400 kton-year exposure for $\sin^2 \theta_{23} = 0.5$ and $\delta_{CP} = 3\pi/2$ as input. Even assuming no prior knowledge on $\sin^2 \theta_{23}$ as done here, a determination of $\sin^2 \theta_{23} = 0.5 \pm 0.1$ can be achieved for maximal mixing at $1\sigma$, with little impact on the $\delta_{CP}$ sensitivity. We leave a detailed analysis to future work, but preliminary results show that the most competitive constraints coming from sub-GeV atmospheric neutrinos are indeed on $\delta_{CP}$. Besides, a similar analysis could be performed for the JUNO experiment [45], as JUNO is also capable of identifying low energy charged particles.

Finally, one could wonder what is the role of neutrino interaction uncertainties in this analysis. We will investigate that in detail in a forthcoming publication, but we would call the attention to the fact that the Short-Baseline Neutrino program [46] at Fermilab and DUNE will explore the sub-GeV region with high statistics. In fact, the DUNE-PRISM concept [47], a movable near detector to probe off-axis neutrinos, could greatly enhance our knowledge of neutrino-argon interactions at sub-GeV scales if the near detector hall allows for an off-axis distance of at least $\sim 25$ meters. This would allow for a pioneering data-driven analysis of $CP$ violation using sub-GeV atmospheric neutrinos in DUNE, fully exploring the unique capabilities of liquid argon time projection chambers.

V. CONCLUSIONS

We have proposed to use the unique capabilities of liquid argon time projection chambers to explore the physics of sub-GeV atmospheric neutrinos. By detecting low energy charged particles, the direction and energy of incoming neutrinos can be inferred, furnishing LArTPCs with a unique opportunity to probe $CP$ violation with low energy atmospheric neutrinos. We have shown, with a detailed simulation, that DUNE’s sensitivity to the $CP$ phase from this atmospheric sample is competitive, possibly ruling out regions of the parameter space beyond the $3\sigma$ level, and will provide a remarkable cross check of the $CP$ phase determination with beam neutrinos. This measurement will have significant consequences in the possible discovery of neutrinos from diffuse supernovae, and in the determination of the neutrino floor in dark matter direct detection experiments, besides opening novel possibilities to probe new physics scenarios. We also highlight the possibility of performing a data-driven analysis using inputs from highly off-axis DUNE-PRISM measurements on neutrino-argon interactions.

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